

Angular Dependence of Jitter Radiation Spectra from Small-Scale Magnetic Turbulence

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ABSTRACT

Jitter radiation is produced by relativistic electrons moving in turbulent small-scale magnetic fields such as those produced by streaming Weibel-type instabilities at collisionless shocks in weakly magnetized media. Here we present a comprehensive study of the dependence of the jitter radiation spectra on the properties of, in general, anisotropic magnetic turbulence. We have obtained that the radiation spectra do reflect, to some extent, properties of the magnetic field spatial distribution, yet the radiation field is anisotropic and sensitive to the viewing direction with respect to the field anisotropy direction. We explore the parameter space of the magnetic field distribution and its effect on the radiation spectrum. Some important results include: the presence of the harder-than-synchrotron segment below the peak frequency at some viewing angles, the presence of the high-frequency power-law tail even for a monoenergetic distribution of electrons, the dependence of the peak frequency on the field correlation length rather than the field strength, the strong correlation of the spectral parameters with the viewing angle. In general, we have found that even relatively minor changes in the magnetic field properties can produce very significant effects upon the jitter radiation spectra. We consider these results to be important for accurate interpretation of prompt gamma-ray burst spectra and possibly other sources.

Subject headings: radiation mechanisms – relativistic plasmas – gamma-rays: bursts

1. Introduction

Radiation from astrophysical or laboratory plasmas depends upon the strength and topology of the magnetic field within the plasma, as well as on the acceleration of plasma particles. For relativistic plasmas, radiative output of an individual particle depends significantly on the particle’s direction of motion, as relativistic beaming modifies the angular distribution of the radiation emitted to lie within a cone of opening angle $1/\gamma$ along the particle’s velocity. Particles in uniform magnetic fields produce synchrotron radiation, which is characterized by the sweep of the radiation cone as relativistic particles orbit or spiral along the magnetic field lines. In turbulent magnetic fields, particles still radiate in a synchrotron regime as long as the scale of the magnetic field variation is longer than the particles’ average Larmor radii. However, if the magnetic field varies on a scale smaller than a Larmor radius, the particles “jitter” through a series of small transverse accelerations without a substantial change of direction. An observer will thus be in a particle’s radiative cone over some length of its accelerated path, until the particle direction is sufficiently deviated that the line of sight is outside the radiative cone. For an isotropic particle distribution, the rates of particles entering and exiting paths within $1/\gamma$ of the line of sight are equal. The resulting radiation will be that of particles “jittering” with small randomized accelerations, and will be reflective of the magnetic field spectra along a line of sight path through the turbulent magnetic field region (Medvedev 2000).

Whereas in synchrotron radiation spectra the frequencies depend upon relativistic beaming and the sweep of the radiation cone, in jitter radiation spectra the frequencies instead depend directly upon the turbulent variations of the magnetic field and relativistic beaming serves only to limit the observed radiation to that emitted by trajectories within $1/\gamma$ of the line of sight (Medvedev 2000). The resulting radiative spectrum differs from that of synchrotron radiation and, in particular, its low-frequency spectral index is not limited

by $s = 1/3$ (i.e., $dW/d\omega \propto \omega^{1/3}$ below the synchrotron peak), referred to in the literature as the synchrotron "line of death" (Preece, et al. 1998).

Gamma-ray bursts (GRBs) are believed to be a natural astrophysical site for the emergence of a strong small-scale turbulent magnetic field. Relativistic shocks generate strong magnetic fields that are random on very small (sub-Larmor) scales via the Wiebel-like (particle streaming) instability (Medvedev & Loeb 1999). In such an instability, a small magnetic field perturbation in the plane transverse to the motions of counter-streaming particles results in the development and growth of filamentary current structures up to a saturation point at which they may persist for some time longer than the dynamical time-scale of the system, creating an extended region of small-scale turbulent magnetic field. This has been studied extensively in numerical PIC simulations (e.g., Silva, et al. (2003); Nishikawa, et al. (2003); Frederiksen et al. (2004); Medvedev, et al. (2005); Spitkovsky (2005); Chang et al. (2008); Spitkovsky (2007)) of both baryonic and pair plasmas. (See Medvedev & Spitkovsky (2009) for a discussion of the applicability and implications of such simulations for GRB physics.) Recent simulations of magnetic reconnection in pair plasmas have also shown the generation of these strong small-scale magnetic fields via the Weibel instability acting on the streams of accelerated particles in the reconnection exhaust funnels (Swisdak, et al. (2008); Zenitani & Hesse (2008)). We speculate that magnetic reconnection events in a GRB may produce electron-positron plasmas in situ, even in initially lepton-poor plasmas. For the resulting regions of strong small-scale magnetic field, synchrotron theory is inapplicable and jitter radiation theory must be considered. Thus, regardless of the particular model of a GRB (e.g., baryonic, leptonic, magnetic, etc.), the jitter radiation mechanism coupled with the relativistic kinematics of the ejected material represents a viable phenomenological model. It has recently been shown that such a model reproduces a number of spectral features of GRB light-curves remarkably well (Medvedev et al. 2009).

Following the approach of Medvedev (2006), we define a geometry where the local filamentation axis (the local axis along which counter-streaming particle motion occurs) lies in the z -direction so that the field perturbation from the shock or reconnection event is amplified by the Weibel instability in the xy -plane (this might be at or upstream from a shock front lying in the xy -plane and propagating in z , or upstream from reconnection exhaust funnels with filaments pointing in z direction and the Weibel fields being perpendicular to it, as suggested by PIC simulations). The resulting amplified magnetic field is randomly oriented in the xy -plane and independently generated at each position in z (shown theoretically in Medvedev & Loeb (1999) and confirmed in PIC simulations such as Nishikawa, et al. (2003), Silva, et al. (2003), and Frederiksen et al. (2004)). The decoupled behavior of the magnetic fields along the filamentation axis (z) and in the plane (xy) transverse to it means that the resulting Fourier spectra are independent of one another and the overall field distribution is highly anisotropic. Qualitatively, the spectrum of the Weibel-generated magnetic field in the direction transverse to the filamentation axis has been shown to rise and then drop at a scale of order the plasma skin depth (Frederiksen et al. 2004). The magnetic field thus has a general spectral form that may be parameterized as

$$f_{xy}(k_{\perp}) = \frac{k_{\perp}^{2\alpha_{\perp}}}{(\kappa_{\perp}^2 + k_{\perp}^2)^{\alpha_{\perp} + \beta_{\perp}}}, \quad (1)$$

where $k_{\perp} = (k_x^2 + k_y^2)^{1/2}$, and $\alpha_{\perp} > 0$ and $\beta_{\perp} > 0$. In this form k_{\perp} refers to the magnetic field wavenumber in the plane transverse to the filamentation axis and κ_{\perp} , α_{\perp} , and β_{\perp} are free parameters controlling the spectral break and the soft and hard spectral indices, respectively. The field along the filamentation axis is in general unknown but we expect it to be of a similar form, with independent free parameters κ_{\parallel} , α_{\parallel} , and β_{\parallel} :

$$f_z(k_{\parallel}) = \frac{k_{\parallel}^{2\alpha_{\parallel}}}{(\kappa_{\parallel}^2 + k_{\parallel}^2)^{\alpha_{\parallel} + \beta_{\parallel}}}, \quad (2)$$

where again $\alpha_{\parallel} > 0$ and $\beta_{\parallel} > 0$. A plot of $f_z(k_{\parallel})$ is shown for a particular choice of

parameters in Figure 1(a) and demonstrates the basic behavior of this function.

The variables k and κ are presumed to be unitless, their units k_0 having been separated into a normalizable coefficient $k_0^{-2\beta}$. We have modified this form from the Medvedev (2006) paper, in which the power in the denominator was simply β and it was required that $\beta > \alpha$. In these spectral forms used here, the asymptotes of the functions given in equations 1-2 are:

$$f(k) = \begin{cases} k^{2\alpha}, & \text{if } k \ll \kappa, \\ k^{-2\beta}, & \text{if } k \gg \kappa. \end{cases} \quad (3)$$

The theory of jitter radiation with the above magnetic field spectra has been utilized by Medvedev (2006) to derive the basic equations describing the jitter spectrum and demonstrate its dependence on the angle θ between the viewing angle and the local filamentation axis. Taking into account our modification of the form of the magnetic field spectra, the analytical work (Medvedev 2006) indicates that the jitter radiation F_ν spectrum should have the following general properties:

1. two breaks, with locations depending on κ_\perp , κ_\parallel , and θ ,
2. a high-energy spectral index β' , where β' approaches β_\parallel as θ goes to 0 and β' approaches β_\perp as θ goes to $\pi/2$,
3. a low-energy spectral index α' , where α' approaches 1 as θ goes to 0 and α' approaches 0 as θ goes to $\pi/2$.

The resulting radiation spectra from numerical calculations for a limited selection of parameter choices and viewing angles were presented in Medvedev (2006). Here we present the results of more extensive numerical calculations illustrating the full range of jitter spectral variation due to viewing angle for a chosen set of field parameters, followed by a more thorough exploration of the magnetic field parameter space and its implications for

the jitter radiation spectrum. Section II of our paper presents the calculated acceleration spectra for a range of viewing angles θ . In Section III we develop the connection between the acceleration and radiation spectra through an analytical treatment of a simple linearized model of the acceleration spectrum. In Section IV we present the radiation spectra calculated for various viewing angles θ and analyze the spectral variation with viewing angle by means of a five-parameter spectral fit. In Section VI we show the effect of variations in the magnetic field parameters upon the radiation spectra. Section VII presents discussion and conclusions.

2. Acceleration Spectra

The equations for calculating the spectrum of a particle’s acceleration due to magnetic field turbulence generated by a relativistic Weibel-type instability were developed in Medvedev (2006), but only calculated in full for a single representative oblique viewing angle of $\theta = 10^\circ$ in between the head-on ($\theta = 0$) and edge-on ($\theta = 90^\circ$) cases. Here we calculate the acceleration spectra for a more complete range of intermediate viewing angles in order to explore the spectral progression with θ .

As derived in Medvedev (2006), the volume-averaged temporal Fourier component of a particle’s acceleration due to the Lorentz force for the static magnetic field generated by the relativistic Weibel instability is

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = (2\pi V)^{-1} \int |\mathbf{w}_{\mathbf{k}}|^2 \delta(\omega' + \mathbf{k} \cdot \mathbf{v}) d\mathbf{k} \quad (4)$$

$$= \frac{C}{2\pi} (1 + \cos^2 \theta) \int f_z(k_{\parallel}) f_{xy}(k_{\perp}) \delta(\omega' + \mathbf{k} \cdot \mathbf{v}) d^2 k_{\perp} dk_{\parallel} \quad (5)$$

where k is the magnetic wavevector, \mathbf{v} is the particle’s velocity vector, and $\mathbf{w}_{\omega'} = \int \mathbf{w} e^{i\omega' t}$. For the case of a shock viewed at an oblique angle θ from the normal of the shock plane (defined as the z -axis), $\mathbf{k} \cdot \mathbf{v} = k_x v \sin \theta + k_z v \cos \theta$, where we have defined the x -axis so that

the velocity vector \mathbf{v} lies in the xz -plane. We can then use the delta function to substitute for either k_x or k_z . This becomes:

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = \frac{C}{2\pi|v \cos \theta|} (1 + \cos^2 \theta) \int f_z \left(\frac{\omega'}{v \cos \theta} + k_x \tan \theta \right) f_{xy} ((k_x^2 + k_y^2)^{1/2}) dk_x dk_y \quad (6)$$

or

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = \frac{C}{2\pi|v \sin \theta|} (1 + \cos^2 \theta) \int f_z(k_z) f_{xy} \left(\left(\left(\frac{\omega'}{v \sin \theta} + k_z \cot \theta \right)^2 + k_y^2 \right)^{1/2} \right) dk_y dk_z \quad (7)$$

where C is an arbitrary normalization constant, proportional to $\langle B^2 \rangle$. The two forms are equivalent; however, as we approach $\theta = 0$ or $\theta = \pi/2$, the calculation is more convenient if one avoids denominators approaching zero by choosing the appropriate form. It should be noted that neither form is valid for the endpoints $\theta = 0$ or $\theta = \pi/2$, which must be treated separately as in Medvedev (2006).

For a single radiating particle, we plug equations (1) and (2) for f_{xy} and f_z into equations (6) and (7) to obtain:

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = \eta_1(\theta) \frac{C}{2\pi v} \int_{-\infty}^{\infty} \int_{-\infty}^{\infty} \frac{(\frac{\omega'}{v \sin \theta} + k_x)^{2\alpha_{\parallel}}}{(\kappa_{\parallel}^2 \cot^2 \theta + (\frac{\omega'}{v \sin \theta} + k_x)^2)^{\alpha_{\parallel} + \beta_{\parallel}}} \frac{(k_x^2 + k_y^2)^{\alpha_{\perp}}}{(\kappa_{\perp}^2 + k_x^2 + k_y^2)^{\beta_{\perp}}} dk_x dk_y \quad (8)$$

where

$$\eta_1(\theta) = \frac{(\tan \theta)^{-2\beta_{\parallel}} (1 + \cos^2 \theta)}{|\cos \theta|} \quad (9)$$

or alternatively

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = \eta_2(\theta) \frac{C}{2\pi v} \int_{-\infty}^{\infty} \int_{-\infty}^{\infty} \frac{k_z^{2\alpha_{\parallel}}}{(\kappa_{\parallel}^2 + k_z^2)^{\alpha_{\parallel} + \beta_{\parallel}}} \frac{((\frac{\omega'}{v \cos \theta} + k_z)^2 + k_y^2 \tan^2 \theta)^{\alpha_{\perp}}}{(\kappa_{\perp}^2 \tan^2 \theta + (\frac{\omega'}{v \cos \theta} + k_z)^2 + k_y^2 \tan^2 \theta)^{\alpha_{\perp} + \beta_{\perp}}} dk_y dk_z \quad (10)$$

where

$$\eta_2(\theta) = \frac{(\cot \theta)^{-2\beta_{\perp}} (1 + \cos^2 \theta)}{|\sin \theta|} \quad (11)$$

This may then be numerically integrated to produce the acceleration spectrum as a function of ω' .

We normalize the wave-vector to a dimensional constant k_0 , hence k is dimensionless, as discussed above. The spectral form remains the same with $\kappa = \kappa' k_0$ and a normalizable factor $k_0^{2\beta}$ in front. The frequencies ω are normalized to ω_0 such that $v = \omega_0/k_0 = 1$. The parameters α_i , β_i , and κ_i in our magnetic field spectra are in general unknown free parameters that depend upon the spatial distribution of the magnetic field in the turbulent region, and the parameters κ_i , corresponding to the dimension correlation lengths of the magnetic field in the turbulent region, will vary even within an astrophysical system and affect the resulting peak energy of the radiation spectrum. For our initial calculations of spectral variations with θ we choose reasonable values of the parameters as follows:

$$\kappa_{\perp} = \kappa_{\parallel} = 10, \quad (12)$$

$$\alpha_{\perp} = \alpha_{\parallel} = 2, \quad (13)$$

$$\beta_{\perp} = \beta_{\parallel} = 1.5; \quad (14)$$

These parameters are varied (jointly and individually) in Section 5 to explore the resulting effect on the calculated radiation spectrum.

We present the results for varying θ in Figure 2(a). Figure 2(b) shows further detail of the critical region around the peak of the acceleration spectra. The graphs show linearly-connected $\log_e - \log_e$ data, with data point intervals of 0.1. We have arbitrarily normalized the spectra so that the low- k part of the $\theta = 2$ spectrum asymptotes at unity. The acceleration spectra are flat for low ω' , then at a certain ω' they turn rapidly into a sloped region, then there is a second transition to a steep decline for high ω' . For θ near 0 the spectra have a clear peak, but as θ increases the peak recedes and eventually disappears altogether. As θ approaches $\pi/2$ a peak again becomes evident but the location of the peak has shifted by about 0.4 from its position for low θ . The transition point for flattening at low ω' appears to move off rapidly to lower ω' as θ approaches 0.

Figure 3 plots the amplitude of the acceleration spectra at our lower calculation

boundary, at the approximate location of the peak for small θ , and at the approximate location of the peak for θ close to $\pi/2$. The crossing of the lines on this graph correspond to spectral transitions as the peak disappears and then eventually reappears at a new location for higher θ . The graph of the slope of the log – log plot in Figure 4 also illustrates the disappearance and reemergence of the peak, but further shows that even for the unpeaked spectra there is a flattening of the slope of the mid-range θ spectra in the region in between the positions of the peaks that appear at higher and lower θ . For unpeaked spectra, there remains a transition region of some extent between the low-frequency and high-frequency power laws; consequently, the unpeaked spectra may still be better fit by division into three power law regions as opposed to two.

An analysis of equations (8) and (10) indicates that ω' functions as a shift of the center of the magnetic spectral form in which we have made the delta-function substitution. Since this ω' factor is always positive, the function's offset always occurs in the direction of negative wavenumber components k_x or k_z . An example of this behavior for the product of two functions such as equations (1) and (2) over a range of offsets is shown in Figure 5. The resulting integral is highly sensitive to the shape of the two functions and the offset between them, which control the resonance like behavior that produces the resulting peak and transition points in the acceleration spectrum. The angle θ plays a role in determining both the width and the offset of one function relative to the other: the $\kappa_{\parallel}^2 \cot^2 \theta$ or $\kappa_{\perp}^2 \tan^2 \theta$ terms in the denominator influence the width of the function under consideration, while the offset is given by $\frac{\omega'}{v \cos \theta}$ or $\frac{\omega'}{v \sin \theta}$. These are linked such that as θ increases, both the width and the offset of the function containing θ increases. The spectral indices α_{\perp} , α_{\parallel} , β_{\perp} , and β_{\parallel} can also influence the width of the functions and hence also affect the location of the transition points in the acceleration spectra. The overall effect of such variations on the resulting jitter radiation spectrum will be explored more in Section 5.

Summarizing the results of this section, we find that the acceleration spectra are generally characterizable to a good approximation by division into three regions: a flat low- k region and two power law regions, as shown in Figure 6. These may altogether be described via one amplitude, two non-zero spectral indices being functions of α s and β s of the field spectra, and two breaks, which depend primarily on the magnetic spectral peaks κ_{\perp} , κ_{\parallel} , and the viewing angle θ . In the next section we develop a simple linearized model of the acceleration spectrum using these five parameters and show that they can be used to analytically predict the behavior of the final jitter radiation spectrum.

3. From Acceleration to Radiation Spectra

The angle-averaged emissivity of a relativistic particle undergoing a series of small transverse accelerations not substantially affecting its overall velocity is as follows (Landau & Lifshitz 1971; Medvedev 2000, and others):

$$\frac{dW}{d\omega} = \frac{e^2\omega}{2\pi c^2} \int_{\omega/2\gamma^2}^{\infty} \frac{|\mathbf{w}_{\omega'}|^2}{\omega'^2} \left(1 - \frac{\omega}{\omega'\gamma^2} + \frac{\omega^2}{2\omega'^2\gamma^4} \right) d\omega' \quad (15)$$

Analysis of the high- ω and low- ω asymptotic behavior of the radiation spectra for shocks viewed head-on and edge-on, as carried out in Medvedev (2006), yields:

$$\begin{aligned} \text{for } \theta = 0, \quad \frac{dW}{d\omega} &\propto \begin{cases} \omega^1 & \text{if } \omega \ll \kappa_{\parallel} v \gamma^2 \\ \omega^{-2\beta_{\parallel}} & \text{if } \omega \gg \kappa_{\parallel} v \gamma^2 \end{cases} \\ \text{for } \theta = \pi/2, \quad \frac{dW}{d\omega} &\propto \begin{cases} \omega^0 & \text{if } \omega \ll \kappa_{\perp} v \gamma^2 \\ \omega^{-2\beta_{\perp}+1} & \text{if } \omega \gg \kappa_{\perp} v \gamma^2 \end{cases} \end{aligned} \quad (16)$$

where θ is the angle between the line of sight and the normal to the shock front (Medvedev 2006) and α_{\perp} , $\alpha_{\parallel} > 1/2$. We thus expect the radiation spectra at oblique viewing angles to vary between these two forms, dominated by the parallel or the transverse spectra as we vary between the two extremes.

The ω dependence contributed by the integral primarily originates in the lower limit $\omega/(2\gamma^2)$ and where it falls on the acceleration spectrum. In section 2, we found that the acceleration spectrum could be simply characterized as three regions of approximately power-law behavior: a flat initial amplitude at low $\log \omega'$ (Region I), a region of positive or negative slope (Region II), and a region with a more steeply negative slope (Region III). Using this to make a simple approximation for our acceleration spectrum in three regions, we can calculate an approximate analytical solution to equation 15.

We define our simple approximation to the acceleration spectrum using five free parameters, all of which may be fit to the spectrum on a $\log_e - \log_e$ plot: $\mathcal{A} = \log a_0$ is the amplitude of the low- ω' limit; $\tau_1 = \log \omega'_1$ is the transition point between the first and second regions; s_1 is the spectral index in the second region; $\tau_2 = \log \omega'_2$ is the transition point between the second and third regions; and $-s_2$ is the spectral index in the third region ($s_2 > 0$). The acceleration spectrum then has the form (shown in Figure 6):

for $\omega' < e^{\tau_1}$:

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = e^{\mathcal{A}} = a_0, \quad (17)$$

for $e^{\tau_1} < \omega' < e^{\tau_2}$:

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = e^{\mathcal{A} - s_1 \tau_1} \omega'^{s_1} = a_0 \left(\frac{\omega'}{\omega'_1} \right)^{s_1}, \quad (18)$$

for $\omega' > e^{\tau_2}$:

$$\langle |\mathbf{w}_{\omega'}|^2 \rangle = e^{\mathcal{A} + s_1 \tau_2 - s_1 \tau_1 + s_2 \tau_2} \omega'^{-s_2} = a_0 \left(\frac{\omega'_2}{\omega'_1} \right)^{s_1} \left(\frac{\omega'}{\omega'_2} \right)^{-s_2} \quad (19)$$

In Figure 7 we show example models for the acceleration spectra at $\theta = 10$ degrees and $\theta = 60$ degrees and their comparison to the original acceleration spectra (as in Figure 2(a)) for a particular choice of fitting rules.

To first order (neglecting the second and third terms in Eq. (15)), the resulting

radiation spectrum is as follows:

for $\omega/2\gamma^2 < e^{x_1}$:

$$\begin{aligned} \left(\frac{dW}{d\omega}\right)_I &= \frac{e^2}{2\pi c^2} e^{\mathcal{A}} \left[(2\gamma^2) - \omega \frac{s_1 e^{-x_1}}{s_1 - 1} + \omega \frac{(s_1 + s_2) e^{-x_1 s_1 + x_2 (s_1 - 1)}}{(s_1 - 1)(s_2 + 1)} \right] \\ &= \frac{a_0 e^2}{2\pi c^2} \left[2\gamma^2 - \frac{s_1}{s_1 - 1} \left(\frac{\omega}{\omega'_1}\right) + \frac{s_1 + s_2}{(s_1 - 1)(s_2 + 1)} \left(\frac{\omega'_2}{\omega'_1}\right)^{s_1} \left(\frac{\omega}{\omega'_2}\right) \right] \end{aligned} \quad (20)$$

for $e^{x_1} < \omega/2\gamma^2 < e^{x_2}$:

$$\begin{aligned} \left(\frac{dW}{d\omega}\right)_{II} &= \frac{e^2}{2\pi c^2} e^{\mathcal{A} - x_1 s_1} \left[\frac{-\omega^{s_1}}{(s_1 - 1)(2\gamma^2)^{s_1 - 1}} + \omega \left(\frac{1}{s_1 - 1} + \frac{1}{s_2 + 1} \right) e^{x_2 (s_1 - 1)} \right] \\ &= \frac{a_0 e^2}{2\pi c^2} \left[-\frac{(2\gamma^2)^{1 - s_1}}{(s_1 - 1)} \left(\frac{\omega}{\omega'_1}\right)^{s_1} + \frac{s_1 + s_2}{(s_1 - 1)(s_2 + 1)} \left(\frac{\omega'_2}{\omega'_1}\right)^{s_1} \frac{\omega}{\omega'_2} \right], \end{aligned} \quad (21)$$

for $\omega/2\gamma^2 > e^{x_2}$:

$$\begin{aligned} \left(\frac{dW}{d\omega}\right)_{III} &= \frac{e^2}{2\pi c^2} e^{\mathcal{A} - x_1 s_1 + x_2 (s_1 + s_2)} \frac{(2\gamma^2)^{1 + s_2}}{|s_2 + 1|} \omega^{-s_2} \\ &= \frac{a_0 e^2}{2\pi c^2} \frac{(2\gamma^2)^{1 + s_2}}{|s_2 + 1|} \left(\frac{\omega'_2}{\omega'_1}\right)^{s_1} \left(\frac{\omega}{\omega'_2}\right)^{-s_2} \end{aligned} \quad (22)$$

The radiative power spectrum (equations (20) - (22)) obtained analytically by our fit-based approximation for the acceleration spectrum agrees with that obtained via full numerical integration in the following section within about 10%. We have chosen our fitting method for the acceleration spectrum to most closely capture the spectral indices; a different choice of fit may allow for a better determination of peak positions.

Calculations of such spectra for parameters fitted to our acceleration spectra at $\theta = 10$ degrees and $\theta = 60$ degrees are shown in Figures 8(a) and 8(b). The region boundaries in these figures indicate that the radiation spectrum cannot be described by a simple linear approximation in the three regions that were defined by the breaks in our acceleration spectrum. The key features (position of spectral peak and spectral breaks) of the radiation spectrum originate in the additional terms in Equations (20) and (21). Consequently, the transition points in our radiation spectrum do not directly correspond to the transition

points in the acceleration spectrum.

The asymptotic behavior of the spectrum at high and low energies can be easily obtained from Eqs. (20) and (22):

$$\begin{aligned} \text{for } \omega/2\gamma^2 < \omega'_1 : \\ \left(\frac{dW}{d\omega} \right)_I &\propto \omega^0, \end{aligned} \tag{23}$$

$$\begin{aligned} \text{for } \omega/2\gamma^2 > \omega'_2 : \\ \left(\frac{dW}{d\omega} \right)_{III} &\propto \omega^{-s_2} \end{aligned} \tag{24}$$

Thus the high and low-energy asymptotic behavior of the radiation spectrum will be identical to the high and low-energy behavior of the acceleration spectrum.

The behavior of the spectrum in the intermediary Region II can be solved for as well. We take the derivative of Equation (21) to solve for the position $P = \log(\omega_p/2\gamma^2)$ of the spectral peak, in cases where it exists:

$$P = \frac{1}{1-s_1} \log \left[\frac{s_1(s_2+1)}{s_1+s_2} \right] + \tau_2 \tag{25}$$

We note that the peak position becomes undefined in the case $s_1 < 0, |s_1| > s_2$, for which Equation (21) is everywhere decreasing. The case $s_1 < 0, |s_1| < s_2$, in which the acceleration spectrum would decline more steeply in Region II than in Region III is impermissible, so we find that the radiation spectrum will be unpeaked whenever the mid-range spectral index s_1 of the acceleration spectrum is negative (i.e., for $s_1 < 0$).

An exploration of the behavior of peak point P relative to the region boundaries shows that for certain values of θ the peak of the Region II function exists, but has crossed the boundary into Region I and consequently does not appear in the resulting radiation spectrum. Thus, while a calculation of P from the results of fitting the acceleration spectrum appears to indicate the re-emergence of a peak in the radiation spectrum as θ

approaches $\pi/2$, this peak falls beyond the Region II lower boundary and is not observed. We note that the first term in Equation (25) is negative for both $0 < s_1 < 1$ and $s_1 > 1$, so the peak is always located below the transition point between Regions II and III.

An analysis of the behavior of Equation(21) below the Region II peak indicates the following behavior:

$$\text{for } \omega'_1 < \omega/2\gamma^2 < \omega'_2 : \left(\frac{dW}{d\omega} \right)_{II} \propto \begin{cases} \omega^1 & \text{if } s_1 > 1, \\ \omega^{s_1} & \text{if } s_1 < 1, \end{cases} \quad (26)$$

Thus, from Equations (23)-(26), it is evident that the spectral indices s_1 and s_2 of the acceleration spectra will generally correspond to spectral indices s_1 and s_2 in two power law regions of the radiation spectra. In the case of the high-frequency spectral index s_2 this correspondence is exact; however, the relation between mid-range (i.e. intermediate-frequency) spectral indices s_1 and s_1 is modified an upper limit of unity on s_1 and also breaks down when the first term in equation 25 is undefined or larger in magnitude than the distance between the acceleration's spectral transition points $\tau_2 - \tau_1$.

The asymptotic form in Equation (23) suggests that we may neglect the second and third terms in Equation (20) and solve for the transition point $T = \log(\omega_t/2\gamma^2)$ at which the dominating term in Region II becomes significant. We find that:

for $s_1 < 1$:

$$\begin{aligned} \left(\frac{\omega_t}{2\gamma^2} \right)^{s_1} &= (1 - s_1)\omega_1'^{s_1} \\ T &= \frac{1}{s_1} \log(1 - s_1) + \tau_1, \end{aligned} \quad (27)$$

for $s_1 > 1$:

$$\begin{aligned} \left(\frac{\omega_t}{2\gamma^2} \right)^{s_1} &= \frac{(s_1 - 1)(s_2 + 1)}{s_1 + s_2} \left(\frac{\omega_1'^{s_1}}{\omega_2'^{(s_1-1)}} \right) \\ T &= \log \left[\frac{(s_1 - 1)(s_2 + 1)}{s_1 + s_2} \right] + s_1\tau_1 + (s_1 - 1)\tau_2 \end{aligned} \quad (28)$$

We have demonstrated that a simple approximation for our acceleration spectrum

allows us to analytically obtain some of the key features of the resulting jitter radiation spectrum, notably that it will have a similar three-region form with $dW/d\omega \propto \omega^0$ for small ω and $dW/d\omega \propto \omega^{-s_2}$ for large ω and a possibly-peaked transition region. Unlike the acceleration spectrum, the intermediary region in the jitter radiation spectrum will have a maximum slope of 1 and may be unpeaked at angles θ at which the acceleration spectrum was peaked. We emphasize that our spectral calculations are all for a single emitting electron, *not* a power-law distribution of electrons. Yet, the power-law photon spectrum emerges at high energies (above the second jitter spectral break ω_2) in this jitter mechanism, in contrast to the synchrotron exponential spectral decay above the synchrotron frequency.

4. Analysis of Radiation Spectra

Now we turn to full numerical calculations of the jitter radiation spectrum, generated by successive numerical integrations of equations (15) and (8) or (10). The results for varying θ are presented in Figure 9(a), with data point intervals of 0.2 on the $\log \omega$ scale. Once again we have normalized the spectra such that the low-energy asymptotic value of the $\theta = 2$ spectrum is unity. The detailed view of the peak region in Figure 9(b) shows results in this region for intervals of 0.05 in $\log(\omega)$. The spectral shapes and trends are, as expected, much like the acceleration spectra but broadened and flattened overall. No peak reemerges in the spectra as θ approaches $\pi/2$. Figure 10 shows the νF_ν spectrum such as is commonly presented for GRBs and used in GRB spectral analysis.

These jitter radiation results show a significant evolution in the spectrum emitted at different viewing angles relative to the main filamentation axis of the magnetic field spectrum. Note that this viewing angle effect in our calculations is entirely due to particles with velocities directed along the line of sight providing the dominant contribution to the radiation emitted to any particular viewing angle. We are neglecting the angular

distribution of the radiation emitted by each particle and using the angle-averaged emissivity for the spectrum emitted in the forward direction by particles with a particular orientation angle relative to the magnetic field filamentation.

Like the acceleration spectrum, the radiation spectrum can be generally described in terms of three regions (two spectral breaks), and an amplitude or slope in each. To conveniently summarize the spectral features and their evolution, we have developed a five-parameter fit which describes the spectral behavior in these three regions, which we designate as Regions R-I, R-II, and R-III to avoid confusion with the Regions I, II, and III as defined earlier for the acceleration spectrum. (Recall that the transition points for the acceleration spectra do not correspond to the apparent transition points in the radiation spectra.) As before, we have chosen our technique to optimize our results for the spectral indices, rather than the spectral transition points.

Since for jitter radiation the soft spectral index varies continuously and approaches 0 for low ω , the results of a simple two-region fit to these spectra would depend significantly upon where the lower bound of the data window falls relative to the peak. Within a fixed data window, a two-region spectral fit would tend to produce an artificial reduction in the soft spectral index for spectra with higher-frequency spectral peaks or breaks. Unfortunately there is no simple way to characterize the behavior of the middle range of the spectrum because of the transition from peaked to unpeaked spectra as θ varies. Even in unpeaked spectra, the extent and curve of the transition region between the flat low ω' part of the spectrum and the strongly negatively-sloped high ω' part of the spectrum varies substantially. Consequently, we have chosen to model the unpeaked spectra still as a three-region spectra rather than solely by its upper and lower asymptotes. We characterize the spectra by defining three lines (each requiring a slope and a reference point) and finding the transition points at which they intersect.

Region R-I (flat, amplitude A): The low-frequency region R-I is flat, with a slope close to zero. To describe this region, we take our initial calculated amplitude to be A , the low-frequency amplitude. For very small θ our lower calculation boundary may be insufficient to capture the initial flat part of the spectrum, since our first spectral transition point approaches $-\infty$ as θ goes to 0.

Region R-II (positive or negative slope s_1): The intermediate-frequency region R-II may have positive slope resulting in a peak or a slight negative slope (of notably less magnitude than the slope in region R-III). Since not all the spectra are peaked, we have chosen to avoid using the peak value for our fit. Instead we define the “drop point” as the region where the second derivative reaches its minimum value. As the place of largest negative change in slope, this coincides well with the “knee” or second break of the function, and is always at slightly higher frequencies than the peak itself. We then find the slope and a reference point in this region by either:

- Method a: for peaked spectra, we take the maximum value of the numerical derivative and its associated data point.
- Method b: for unpeaked spectra, we take the average value of the numerical derivative in the region between the drop point and the “deviation point” where the spectrum first drops below $A - 0.01$ (this corresponds to a deviation of about 1% from its original value). Our reference point is the data point halfway between or next highest to halfway between the deviation and drop points.

Region R-III (negative slope, defined as $-s_2$): The high-energy frequency region R-III has a large negative slope compared to the rest of the function. This slope is still changing over the region close to the second spectral break that we are considering, so we determine a representative slope by calculating the slope of a line between the drop point, which is

the minimum of the numerical second derivative, and the higher frequency point that is the absolute minimum of the numerical second derivative (the data point closest to where the second derivative crosses zero in this region). These points are well-defined for all our radiation spectra as long as the calculation boundary extends a couple orders of magnitude in e above the drop point. Either point may be used as a reference point in this region.

The first spectral transition point τ_1 is obtained by solving for the intersection of the lines defined in Regions R-I and R-II, and the second spectral transition point τ_2 is obtained by solving for the intersection of the lines in Regions R-II and R-III.

We have chosen to work with the F_ν spectrum because of the convenience of its distinctive flat (spectral index of 0) initial amplitude at very low frequencies, but it is easy enough to translate F_ν spectral features into spectral features of the νF_ν spectrum or the photon spectrum $N(E)$, as the spectral indices will simply be increased or decreased by 1 and the transition points between the power law regions will roughly coincide, with a slight shift based on normalization. In terms of the Band function fit commonly used for GRB spectra (Band et al. 1993), the relation between the high-energy spectral indices is $\beta_{Band} = s_2 - 1$. The relation between the low-energy spectral indices is complicated by the fact that the Band function is a two-region fit and not sensitive to multiple spectral indices below the spectral peak; consequently α_{Band} will range between -1 and $s_1 - 1$ depending on where the data fitting window falls relative to our first spectral break τ_1 . The νF_ν peak energy, which is E_p in the Band function, will correspond to a slope of -1 in the F_ν spectrum, and will lie roughly in the vicinity of the second spectral break τ_2 in our fit.

Figures 12 - 14 show spectral fit results obtained using our above technique on the radiation spectra. We have also applied the same technique to the acceleration spectra presented in section 2 and plotted them for comparison. In addition to data resolution effects, some discontinuity in fitting the peaked vs. unpeaked form of the spectrum is

unavoidable and is reflected in our results.

Figures 12 and 14 indicate that the amplitude A and the high-frequency spectral index s_2 are close in both their values and their evolution with θ for the two types of spectra. The mid-range spectral index s_1 varies similarly in Fig. 13 for both spectra, but appears to approach different asymptotic values as it approaches $\theta = 0$ and $\theta = \pi/2$, clearly showing the expected $s_1 < 1$ limiting behavior. The second spectral break τ_2 shows that the radiation transition point tends to be about $1/2$ a power of e lower than the second spectral behavior for the acceleration case, but shows similar evolution with θ in both cases.

Figure 17 shows the angular dependence of the spectral peak in both our acceleration and radiation F_ν spectra, the νF_ν spectral peak (peak data point in Figure 10, and the drop point, which we have defined as the minimum in the numerical second derivative. We clearly see the usefulness of the drop point in tracking the spectral behavior across the full range of θ , and that it closely tracks the behavior of the νF_ν spectral peak (E_p). Both Figures 17 and 18 clearly show the re-emergence of peaked acceleration spectra as θ approaches $\pi/2$ and the lack of peaked radiation spectra for similar values of θ .

5. Spectral Features and Exploration of the Spectral Parameter Space

We have explored the influence of changes in the magnetic field spectral parameters on the acceleration experienced by the particle and hence its resulting radiative profile. Sections 2 and 4 presented the acceleration and radiation spectra calculated from magnetic field spectra of the form given in Equations (1) and (2) with our original choice of parameters $\alpha = \alpha_\perp = \alpha_\parallel = 2.0$, $\beta = \beta_\perp = \beta_\parallel = 1.5$, $\kappa = \kappa_\perp = \kappa_\parallel = 10$. In this section, we present the results of varying these parameters. We vary the joint parameters α , β , and κ in the parallel and perpendicular magnetic field spectra. We also vary the parameters α_\parallel ,

α_\perp , β_\parallel , and β_\perp individually. Finally, we vary the ratio $K = \kappa_\perp/\kappa_\parallel$. For each variation of the initial parameters, we have calculated the radiation spectrum for three representative angles at $\theta = 10^\circ$, 45° , and 80° . The results are presented according to their impact on the characteristics of the radiation spectrum as developed in section 4, namely the initial amplitude A , the spectral breaks τ_1 and τ_2 , and the spectral indices s_1 and s_2 . We also present the results for the peak "strength", the height of the spectral peak above the initial amplitude.

The spectra are divided by $\langle B^2 \rangle \propto \int f(\kappa_\parallel) f(\kappa_\perp) d^3k \propto k_0^{-2(\beta_\parallel + \beta_\perp)}$ to appropriately normalize the amplitudes relative to one another, but in all cases we have arbitrarily normalized the final spectra such that the low-energy asymptotic value of the $\theta = 10^\circ$ spectrum with our original choice of parameters is unity (zero on the logarithmic scale).

As in Section 4, the initial, low-frequency amplitude A is the first calculated value of the angle-averaged radiative power emitted per frequency $dW/d\omega$. This value is generally a good approximation for the asymptotic value of the function as it approaches lower ω , though it may deviate somewhat from this value for θ approaching 0, as the spectra becomes sloped rather than flattened at our lower calculation boundary in ω .

Among the resulting figures 19(a)-20, variations in the magnetic field parameters κ_i produce the largest effect on the low-frequency amplitude, causing changes of about 4 orders of magnitude in e when varied individually via the ratio $K = \kappa_\perp/\kappa_\parallel$, and up to 7 orders of magnitude in e when varied together as $\kappa = \kappa_\parallel = \kappa_\perp$. Variation with changes in the magnetic field spectral indices α_i and β_i are small in comparison, on the scale of about 1-2 orders of magnitude. The amplitude increases with increasing K for $\theta = 10^\circ$ and generally decreases with increasing K for $\theta = 80^\circ$; thus, it increases when κ_\perp dominates at small θ and when κ_\parallel dominates at large θ .

The mid-range spectral index in peaked spectra is the maximum slope below the

peak, as determined by taking the numerical first derivative of our calculated values. For unpeaked spectra, we find the spectral index as the average slope between the point at which the graph deviates by more than 0.01 from the initial amplitude A and the drop point at which the numerical second derivative reaches a minimum.

Figures 21(a)-21(c) present the effect of magnetic field parameter variations on s_1 . As can be seen in Figure 22, the mid-range spectral index is strongly affected by variations in both the parameters β_i , especially for $\theta = 10^\circ$, and in the relative strength of the κ_i values. The peaked $\theta = 10^\circ$ spectra is notably more sensitive to magnetic field variations than the unpeaked 45° and 80° spectra. The ratio $K = \kappa_\perp/\kappa_\parallel$ has the largest influence at all three representative viewing angles, with s_1 increasing as $\kappa_\parallel > \kappa_\perp$). We note that even at θ approaching $\pi/2$, we obtain a positive slope (and hence a peaked spectrum) for $K = 1/10$.

For both peaked and unpeaked spectra, the high-frequency spectral index $-s_2$ is determined by taking the slope between the drop point and the absolute minimum of the second derivative above the peak (i.e., the closest data point to where the second derivative crosses zero). Figures 23(a) - 23(c) show the effects of variations in the magnetic field parameters on $-s_2$. We find that as expected analytically, this spectral index is primarily influenced by the magnetic field parameters β_i . In particular, s_2 is most strongly influenced by β_\perp , the high-wavenumber spectral index of the magnetic field spectrum transverse to the current filamentation. As seen in figure 23(b), β_\parallel affects s_2 only at small angles θ , and its influence even then is less than that of varying β_\perp . The apparently strong influence of κ is largely an artificial effect as the κ parameter's strong shifting of the function (as indicated in our analysis of the spectral breaks below) towards higher frequencies interferes with the calculation of s_2 by shifting the absolute minimum of the second derivative outside our calculation boundaries. This causes an artificial reduction in the steepness of the slope for $\kappa = 100$, as evident also in Figure 23(c).

We calculate the first spectral break (i.e. transition point) as the intersection between the line $\log \langle |\mathbf{w}_{\omega'}|^2 \rangle = A$ and the line of slope s_1 through the point of maximum positive slope for peaked spectra or through the data point in the middle of the range over which we averaged to obtain slope s_1 for unpeaked spectra. (If the middle of the range does not fall on a data point, we take the next larger data point.) We find, as shown in figures 25(a)-26, that the first spectral break is strongly influenced by changes in α and κ or the κ -ratio K . The break position shifts to higher frequency by about an order of magnitude in e as we increase α from 1 to 10 jointly in the parallel and perpendicular magnetic field spectra. In varying the α_i separately we see that α_{\parallel} has a larger influence at $\theta = 10^\circ$ and α_{\perp} has a larger influence at $\theta = 80^\circ$. Increasing the κ jointly by powers of 10 results in shifting the first spectral break to higher frequencies by roughly 4 orders of magnitude in e . Varying the κ parameters relative to one another results in a similarly strong shift, towards higher frequencies for $\kappa_{\perp} > \kappa_{\parallel}$.

The second spectral break τ_2 is calculated as the intersection between a line of slope s_1 through the point of maximum positive slope (for peaked spectra) or through the mid-point of the averaging region (for unpeaked spectra), and the line of slope $-s_2$ through the "drop point" at which the second derivative reaches a minimum (i.e., the largest negative change in the slope). Our results (in figures 27(a)-28(c)) indicate that the second transition point is most strongly influenced by the κ_i varied jointly or via the ratio K . The low-wavenumber magnetic field spectral index also demonstrates a fairly strong influence, with larger α shifting τ_2 to higher frequencies. A comparison of the influence of α on the two spectral break points (as seen in Figures 25(a) and 27(a)) indicates a very similar shift in both break points; thus increasing α shifts the entire spectrum towards higher frequencies.

Figures 29(a)-29(c) show the variation in the peak strength (which we have defined as the height of the spectral peak above the low-frequency amplitude A with changes in the

magnetic field parameters. For $\theta = 10^\circ$, the only of our three representative angles that has peaked spectrum for $K = 1$, peak strength increases with increasing α and β . Individually, increasing β_\perp has the largest effect in increasing the peak strength, while increasing α_\perp lowers it. Similarly, increasing α_\parallel increases the peak strength while increasing β_\parallel lowers it. The largest effect overall is produced by variation of the ratio K between the perpendicular and parallel field parameters κ_i . For $\kappa_\parallel > \kappa_\perp$ (i.e. $K < 1$), the peak strength appears to persist to higher angles θ , while for $\kappa_\perp > \kappa_\parallel$ the peak can be small or non-existent even at $\theta = 10^\circ$. Thus the ratio between κ_\perp and κ_\parallel , the respective peaks of the magnetic field perpendicular and parallel spectra, strongly influences the progression of the spectral evolution between its $\theta = 0$ and $\theta = \pi/2$ limiting values, as expected from our earlier analysis in Section 2.

We have seen that relatively minor changes in the magnetic field spectra can produce very significant effects upon the jitter radiation spectra, particularly in the appearance of the spectral peak or break region. Furthermore, while we have included in this section only spectra from a few representative viewing angles θ , the angular dependence demonstrated suggests that the connection of such features to the transverse or parallel magnetic field spectra can be tested by observing their variation with viewing angle.

6. Conclusions

We have calculated the angle-averaged power spectra of jitter radiation emitted by *a single relativistic electron* undergoing small Lorentz-force accelerations transverse to its overall velocity. Note that the obtained spectra are equivalent to the ensemble-averaged spectra per one electron from a collection of monoenergetic relativistic electrons. The radiation spectra are calculated using a smoothly connected broken power-law model of a magnetic field mimicking the structure of magnetic fields generated by the Weibel

instability. The shapes of the resulting jitter radiation spectra are shown to depend on the magnetic field spatial spectrum and to vary with the angle θ of the electron velocity (being also the line of sight) with respect to the direction of the field anisotropy (z -axis). The effect of varying parameters in the magnetic field spectra has been explored and indicates that the jitter radiation spectral features, such as the strength of the spectral peak or the extent of a sloped transition region, are quite sensitive to the parameters controlling the magnetic field spectra.

Despite the high sensitivity of the jitter radiation spectra to the magnetic field spatial spectrum or, in general, the field correlation tensor, $K_{ij}(\mathbf{k}) = B_{\mathbf{k}}^i B_{\mathbf{k}}^{*j}$ (Medvedev 2006), one can draw some fairly robust conclusions. When the parallel and perpendicular magnetic field spectra are similar, one has just four essential parameters, their low- k and high- k spectral slopes, $\alpha > 1/2$ and $\beta > 0$, the peak representing a typical correlation length κ , and the viewing angle of the line of sight with respect to the magnetic filament direction, θ . The power (i.e., F_ν) spectrum produced by monoenergetic electrons moving towards the observer with the Lorentz factor γ , in general, has three power-law segments: a flat low-energy part, an intermediate-energy region which rises or slightly falls with a slope of less than unity (the exact value depending on θ), and the more steeply falling off part with the slope being between -2β and $-2\beta + 1$, again, depending on θ .

The shape of the spectrum changes significantly with the angle θ between the radiating particle's velocity and the axis of the current filamentation generated by the counterstreaming Weibel instability. As $\theta \rightarrow 0$, the low-frequency spectral break τ_1 approaches $-\infty$ and the maximum spectral slope (mid-range spectral index s_1 approaches the value of 1 (the trend of our results agreeing well with the $\theta = 0$ case in Medvedev (2006)). As θ increases, the spectral peak weakens as s_1 decreases and τ_1 shifts towards the peak region. The disappearance of the spectral peak at some particular θ appears to be a

result of both these spectral changes, and there is an extended transition region between the low-energy and high-energy power law trends. Consequently, we find that both the peaked and unpeaked spectra are well described by a three-region fit. Two-region fits are likely to miss out on the variation in the spectral slope below the peak at small θ ; consequently the resulting low-energy spectral index will be extremely sensitive to where the peak falls relative to the lower bound of a measured spectral window. This will be true even if the low-energy spectral index is taken at a common energy, as in the “effective” low-energy spectral index α_{eff} commonly taken as the tangential slope of the logarithmic spectrum at 25 keV.

In comparing the radiation spectra for the full range of θ , we have found that the “drop point”, which we determined as the minimum of the numerical second derivative of the logarithmic data (i.e. the largest negative change in the spectral slope) serves as a good common reference point for both peaked and unpeaked spectra; in addition, the “drop point” in the radiation power spectrum evolves with the angle θ much like the νF_ν spectral peak energy E_p .

In section 3, we developed in detail the relation between the radiation spectrum and the underlying Fourier spectrum of the particle’s acceleration. In particular, we find that the radiation spectrum has much the same shape as the acceleration spectrum but that the apparent transition points in the two spectra do not simply coincide for most angles of θ . Furthermore, although the acceleration spectrum sees the re-emergence of a spectral peak for $\theta \rightarrow \pi/2$, the radiation spectrum does not. We have also demonstrated that a simple fit to the acceleration spectrum allows for the generation of a model radiation spectrum which approximates the realistic one with 10% accuracy.

We have found that variations in the magnetic field spectral parameters influence the final radiation spectrum by controlling the width and peak-positions of the functions within

the integrand and the extent to which this directly modifies the effect of the offset, which is proportional to ω' . If we consider the general progression of the radiation spectra from being strongly peaked at small θ to unpeaked at θ near $\pi/2$, the trends shown here indicate that the speed of the progression of the spectral shape between the two extremes is dependent on the relative strengths of the parameters in the magnetic field spectra transverse and parallel to the shock front.

We have confirmed that the jitter radiation high-energy spectral index is determined primarily by the high- k magnetic field spectral index β , which otherwise has little influence on the spectrum. The low- k magnetic field spectral index α is shown to have a significant influence on the low-energy and mid-range portions of the radiation spectrum when varied jointly in both magnetic field spectrum, although this influence is substantially reduced when only one α_i is varied.

The parameters κ_{\perp} and κ_{\parallel} represent the dimensionless correlation lengths of the magnetic field distribution in the direction along the Weibel current filaments (and the direction of shock propagation in the case of a GRB) and in the perpendicular plane (parallel to the shock plane for a GRB). We find that increasing κ_{\perp} and κ_{\parallel} jointly shifts the entire spectrum to higher energies with relatively little effect on the spectral shape. Thus, as expected, the location of the spectral peak and break energies (and the corresponding peak energy E_p of the νF_{ν} spectrum) are determined primarily by the correlation length of the magnetic field turbulence. The progression of the spectral shape between the head-on and edge-on cases is sensitive to the variation of the κ parameter in one function relative to the other, such that for a particular viewing angle θ either peaked or unpeaked spectra can be attained via modification of the κ ratio K . In the extreme that κ_{\parallel} is 2 orders of magnitude larger than κ_{\perp} we recover a peaked spectra for the angles as high as $\theta \sim 80$ degrees. It is also notable that the spectral peak and transition points undergo relatively

little horizontal shift as K varies when $\theta = 10$ degrees, but shift quite dramatically (3-4 orders of magnitude) during this variation for $\theta = 80$ degrees.

We shall summarize the most notable properties of our jitter radiation results with particular significance to interpreting radiation spectra from astrophysical sources. First, the jitter radiation spectra are significantly harder than synchrotron spectra in the region just below the spectral peak. This may be a significant mechanism in astrophysical sources such as gamma-ray bursts where a substantial population is seen to violate the synchrotron limit. Second, both the maximum slope in the region below the peak and the extent of this sloped region (between the low-frequency spectral break and the peak) are strongly influenced by θ , with the largest slope and largest extent of the sloped region at small θ . The angle θ_{np} at which the peak disappears is determined primarily by the ratio between the magnetic field correlation lengths perpendicular to and along the filamentation axis. Third, the position of the spectral peak represents the characteristic correlation length of the magnetic field, which in case of the Weibel turbulence depends on the density as $n^{1/2}$. This is at odds with the synchrotron radiation in which the spectral peak measures the magnetic field strength. This result is important for accurate interpretation of the observed spectra as well. Fourth, the high-energy part of the spectrum is represented by a power-law, even though the electrons are monoenergetic. Thus, no power-law distribution of Fermi-accelerated electrons is required to produce the observed power-law spectra in prompt GRBs. This has important implication for the interpretation of the observed data, provided the electrons are radiating in the jitter regime. Fifth, the angular dependence exhibited, combined with relativistic kinematics of a curved shock front can explain certain puzzling features of the GRB prompt spectral variability (Medvedev et al. 2009). The sensitivity of the jitter spectra to the magnetic field anisotropy makes it a possible tool for diagnostics in sites of small-scale magnetic field turbulence and also a basis for analysis of astrophysical sources where the magnetic field orientations relative to the direction of

observation are changing over time.

Although we often appeal to GRBs as sites where jitter radiation can likely be produced, we cannot exclude other astrophysical objects, e.g. jets in active galactic nuclei, early supernova shocks and other violent sources, provided that small-scale magnetic fields may be produced and maintained in them. At last but not least, one can use jitter radiation as diagnostic tool in laser-plasma interaction experiments, e.g., Hercules (Huntington et al. 2008; Reynolds, et al. 2007), aimed at studies of Weibel turbulence and conditions in GRBs within the Laboratory Astrophysics and High-Energy-Density Physics programs.

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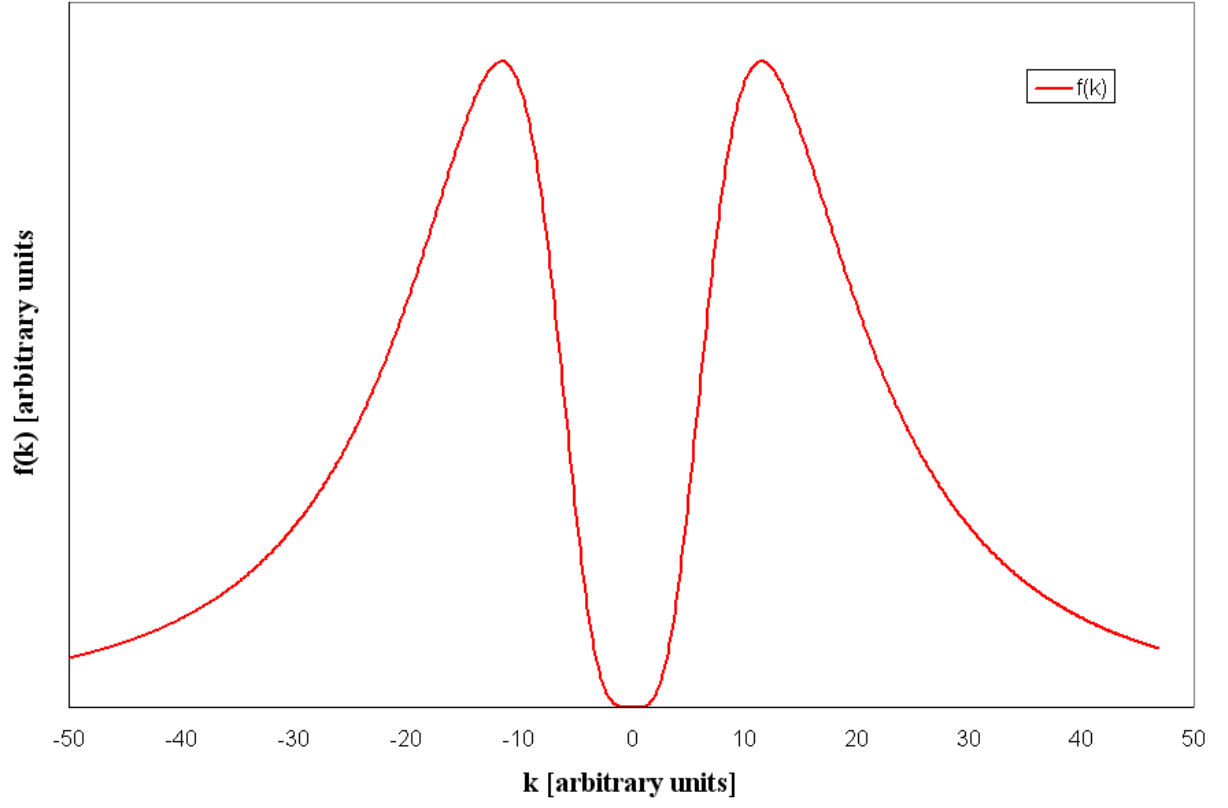


Fig. 1(a).— A plot of the general form of the spectrum $f(k)$ used in our paper for the magnetic field spectrum in the transverse and parallel directions.

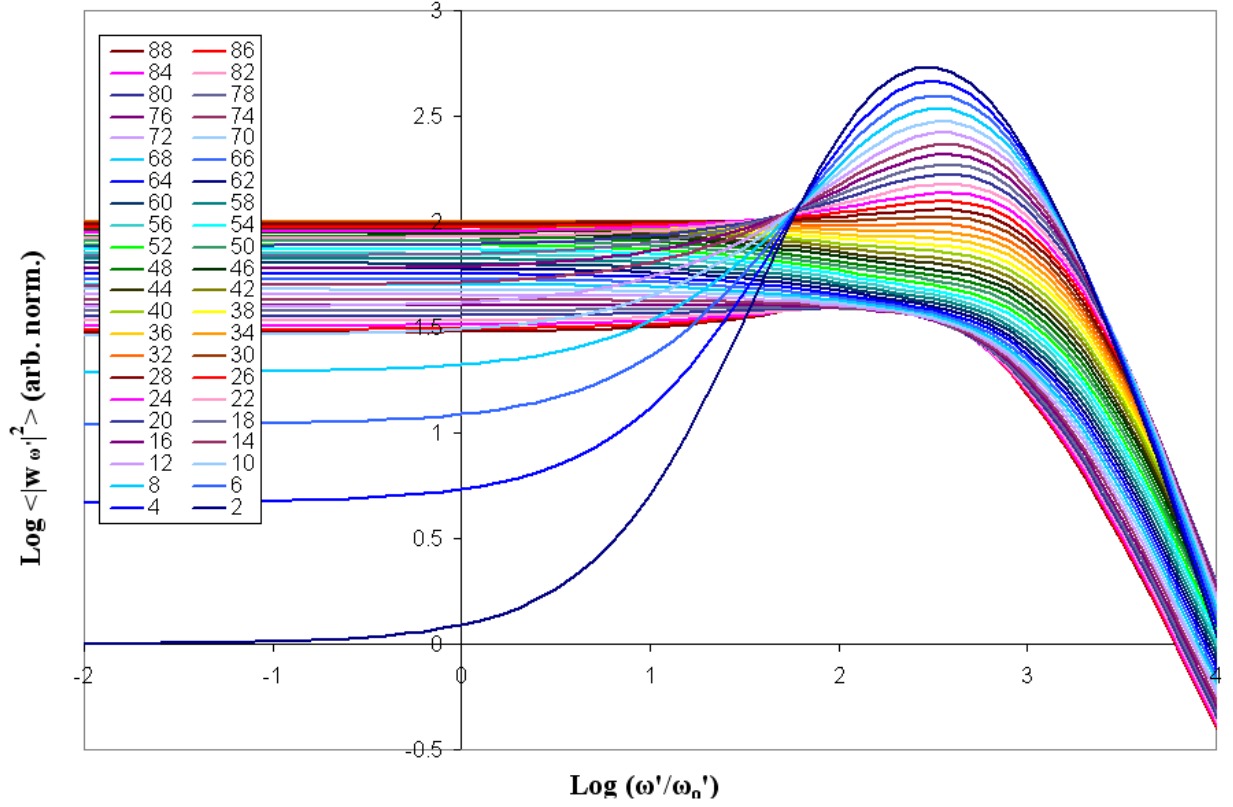


Fig. 2(a).— The acceleration spectra $\langle |\mathbf{w}_{\omega'}|^2 \rangle$ for θ ranging from 2 through 88 degrees. The spectra is numerically calculated for a step size of 0.1 in $\log(\omega')$ and is normalized via division by the 3-dimensional integral over the magnetic field spectra.

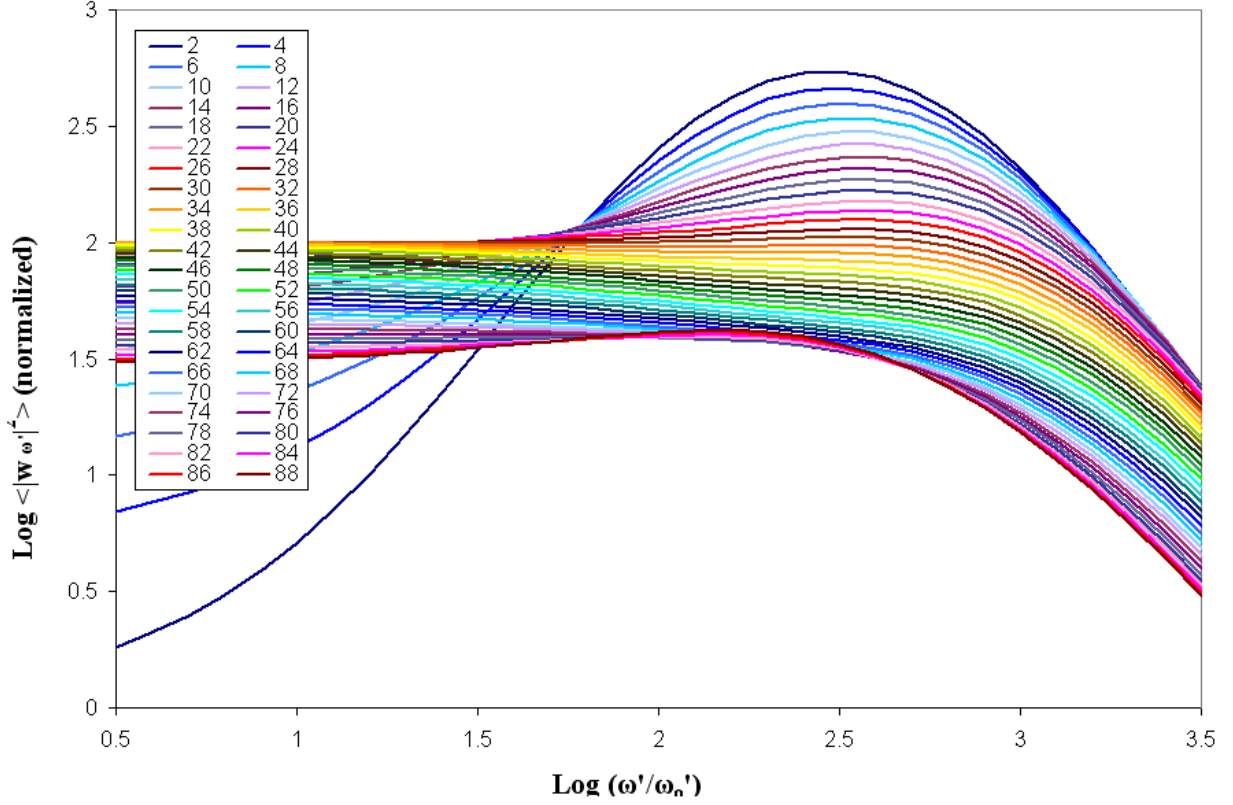


Fig. 2(b).— View showing detail of peak region of acceleration spectra $\langle |\mathbf{w}_{\omega'}|^2 \rangle$ for θ shown in Figure 2(a). One can see the disappearance of the peak and flattening of the spectra for mid-range θ , followed by its reappearance at slightly lower $\text{log}(\omega')$ at θ of about 76 degrees.

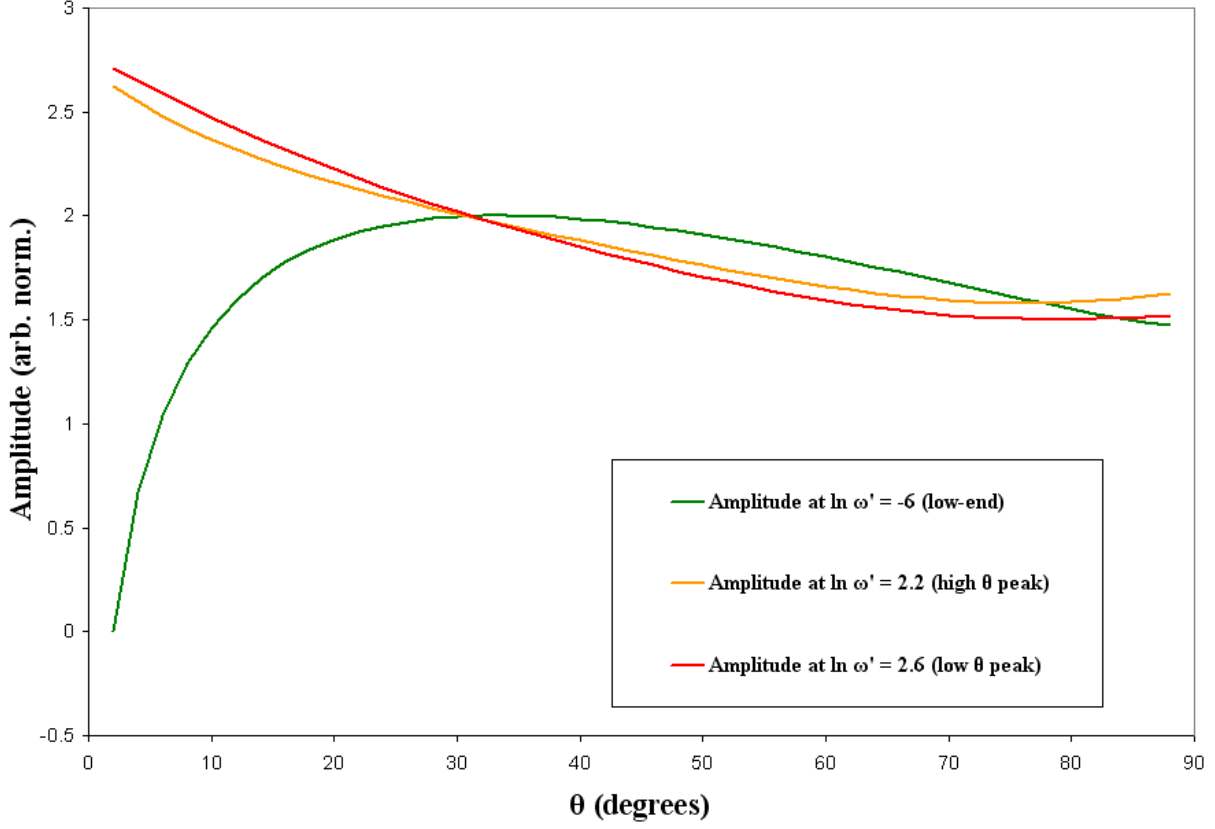


Fig. 3.— Plot of the spectral amplitudes taken at the low- ω' end of the calculated spectra, and at the approximate locations of the spectral peak for low- θ and high- θ . The dominance of one amplitude over the other illustrates the transition of the spectra from peaked to unpeaked as θ progresses from 0 to $\pi/2$.

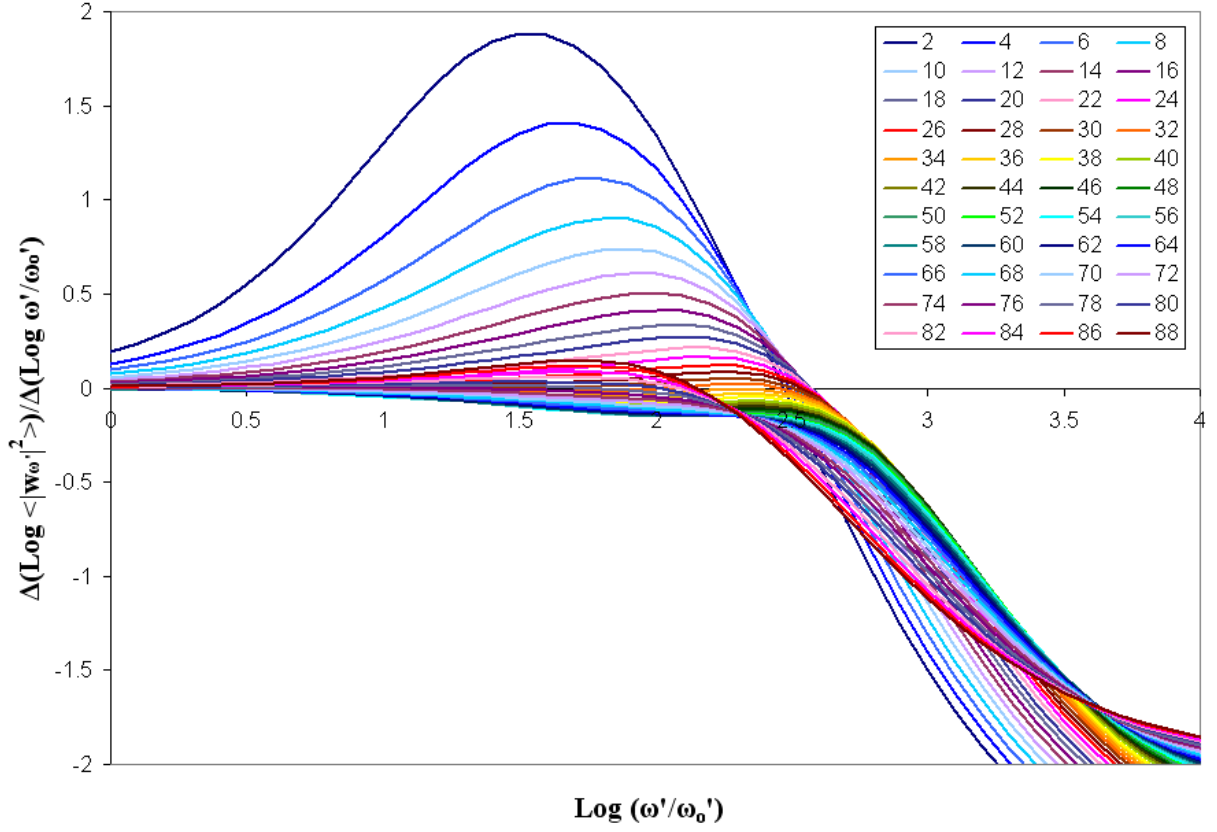


Fig. 4.— The slope of the spectra in the log – log plot in Figure 2(a). Note that even for unpeaked spectra there is a flattening of the spectral slope and in some cases a local maximum (around $\log(\omega'/\omega_o') = 2.5$).

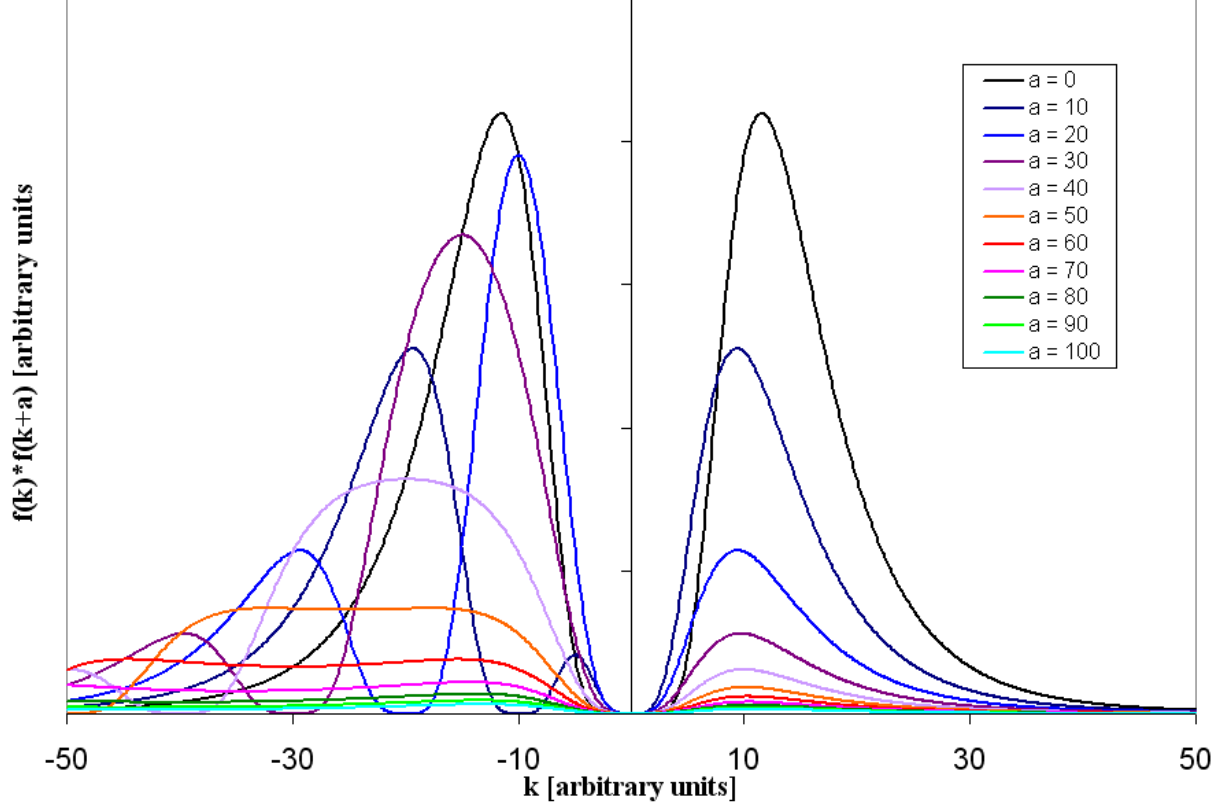


Fig. 5.— Plots of the product $f(k)f(k+a)$ of two functions of the form of our magnetic field spectra (Equations (1) and (2)), illustrating the effect of the offset a . As described in section 2, the acceleration spectrum is the integral of a product of such functions, with an offset controlled by the acceleration Fourier frequency ω' .

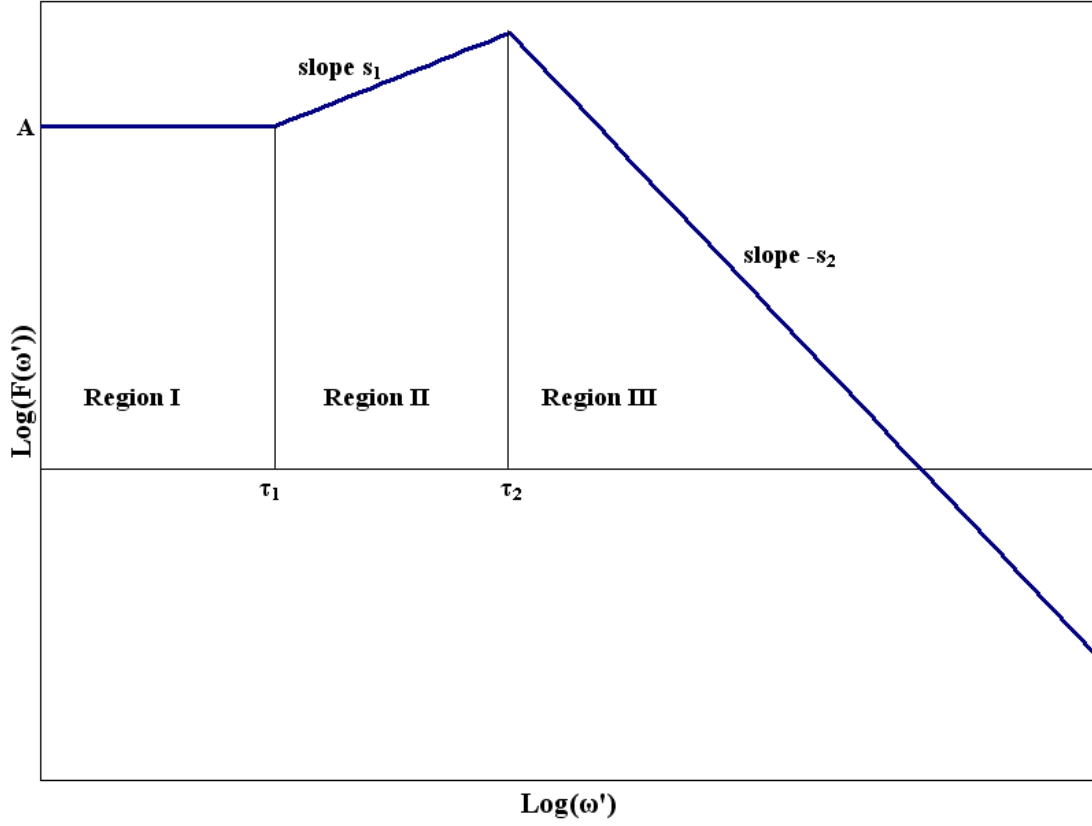


Fig. 6.— Here we show a model acceleration spectrum, as described in Equations (17) - (19). The spectrum is flat in Region I ($\log(\omega') < \tau_1$), then becomes a power law ω'^{s_1} in Region II ($\tau_1 < \log(\omega') < \tau_2$), and a power law ω'^{-s_2} in Region III ($\log(\omega') > \tau_2$).

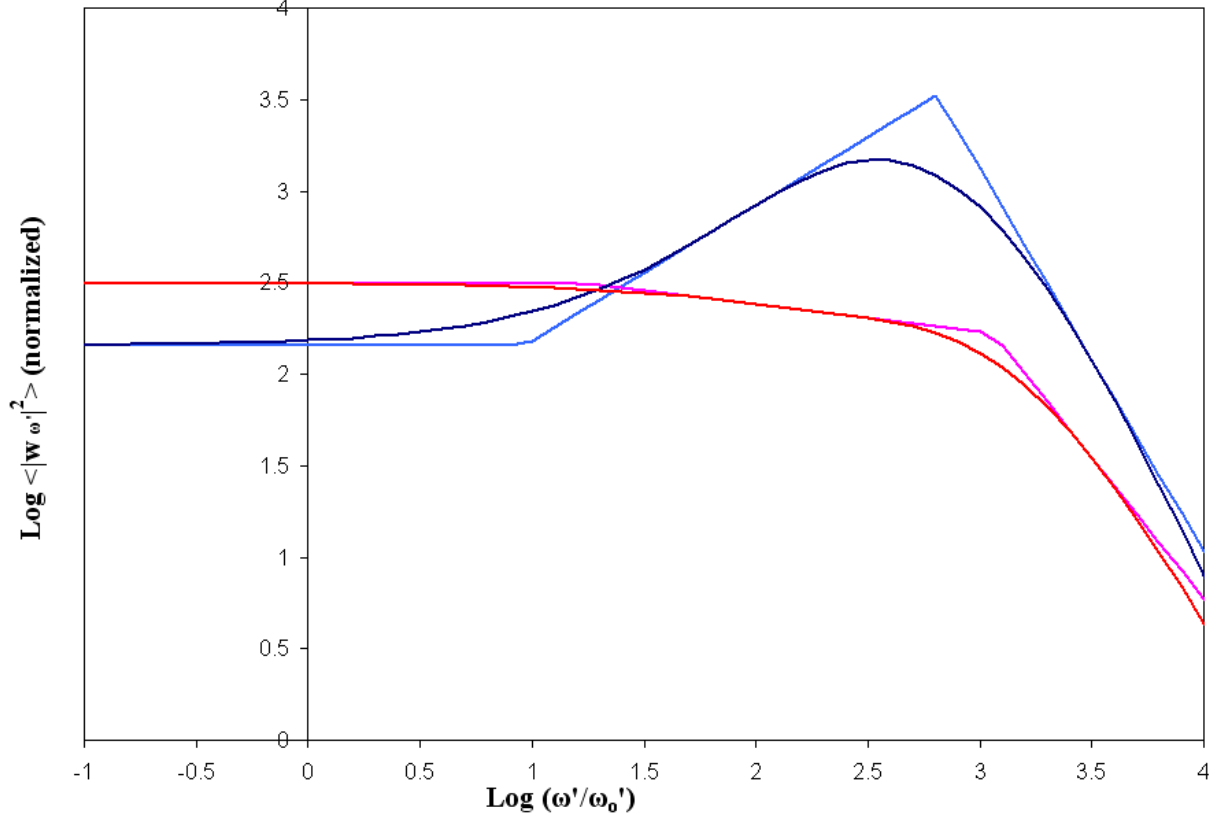


Fig. 7.— Here we show a comparison of our original calculated spectra (as presented in Figure 2(a)) for $\theta=10$ (peaked) and $\theta=60$ (unpeaked), and our model acceleration spectra as described in Equations 17-19. In Region I, we extrapolate the initial calculated amplitude. In Region II, we fit the calculated spectra using the maximum slope in the region and fitting through the point of maximum slope. In Region III we have done a simple fit using the slope and position at point $\omega=4.6$. Our choice of fit overestimates the peak of the calculated acceleration spectra.

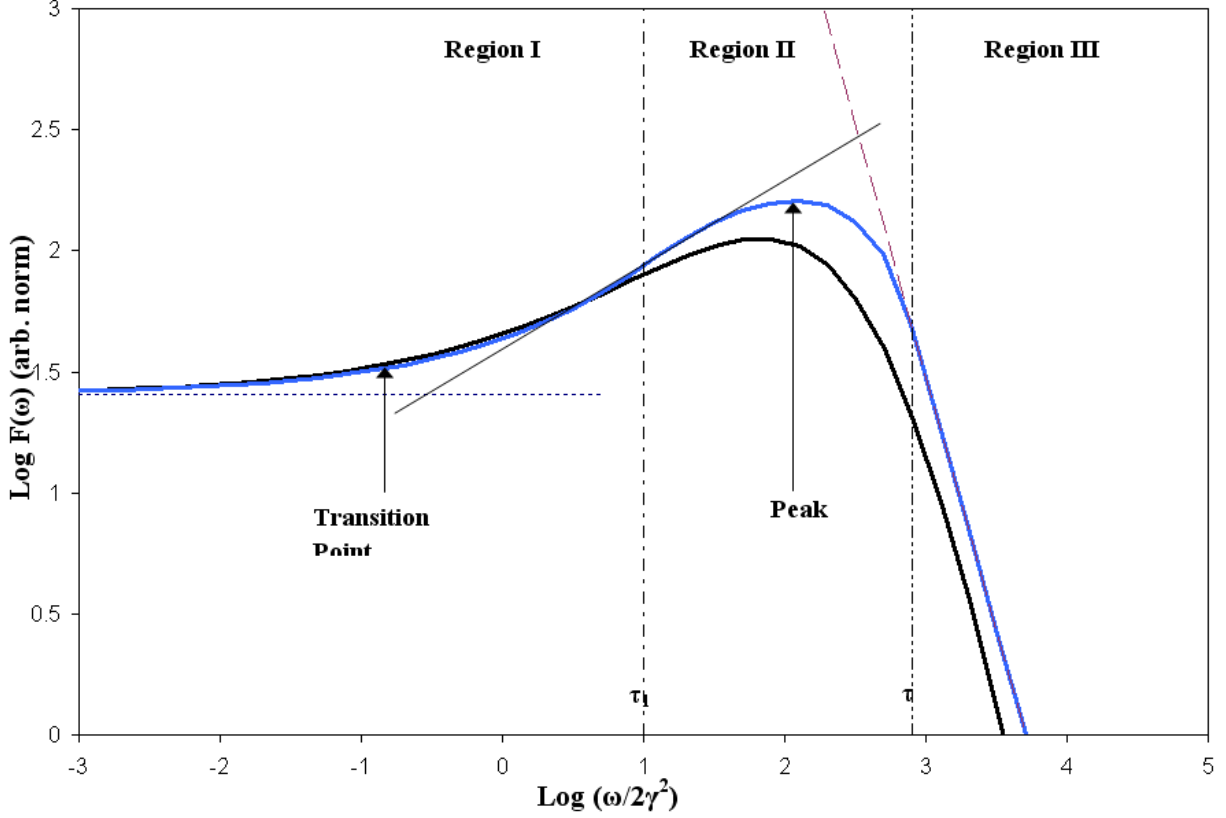


Fig. 8(a).— The radiation spectrum for $\theta=10$ degrees as obtained analytically using equations 20 - 22 and the fit values for the model acceleration spectra as shown in Figure 6. The result calculated (black) via full double numerical integration (as in Figure 9(a)) is shown here for comparison. The spectral transition point defined by equation 27 and the peak defined by equation 25 are marked with arrows. We note that the boundaries (dotted lines) between Regions I, II, and III as defined by our model acceleration spectra do not clearly correspond to transition points in the resulting radiation spectra. We have extrapolated slopes for Regions I, II, and III to demonstrate the variation. The model again slightly overestimates the peak energy, but matches the spectral shape well and agrees with the integrated spectra within about 10%.

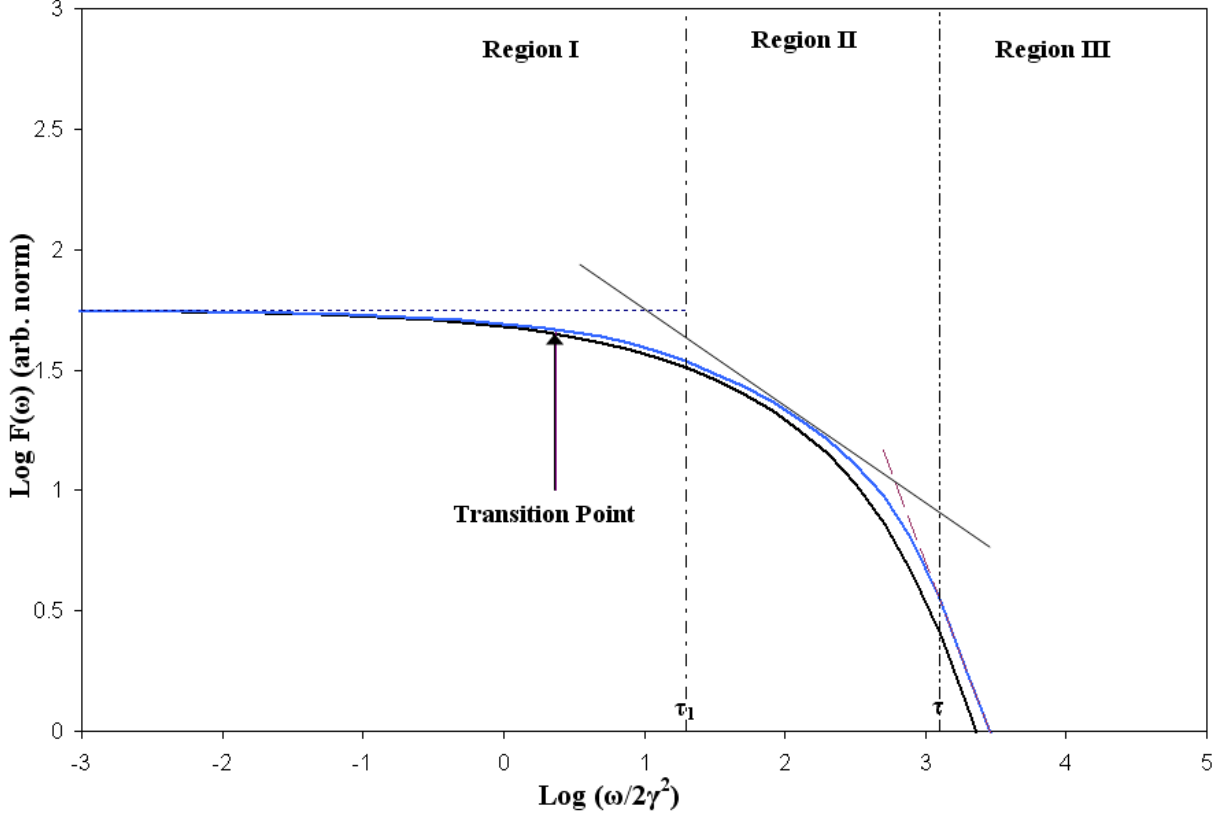


Fig. 8(b).— The radiation spectrum for $\theta=60$ degrees as obtained analytically from our simplified model acceleration spectrum (as shown in Figure 6) versus the full double numerical integration result (black). As in the previous figure, we have extrapolated the slope in Regions I, II, and III. We again obtain a good agreement with the overall spectral shape, but now with an unpeaked form appropriate to this range of theta. The spectral transition point defined by equation 27 is marked with an arrow.

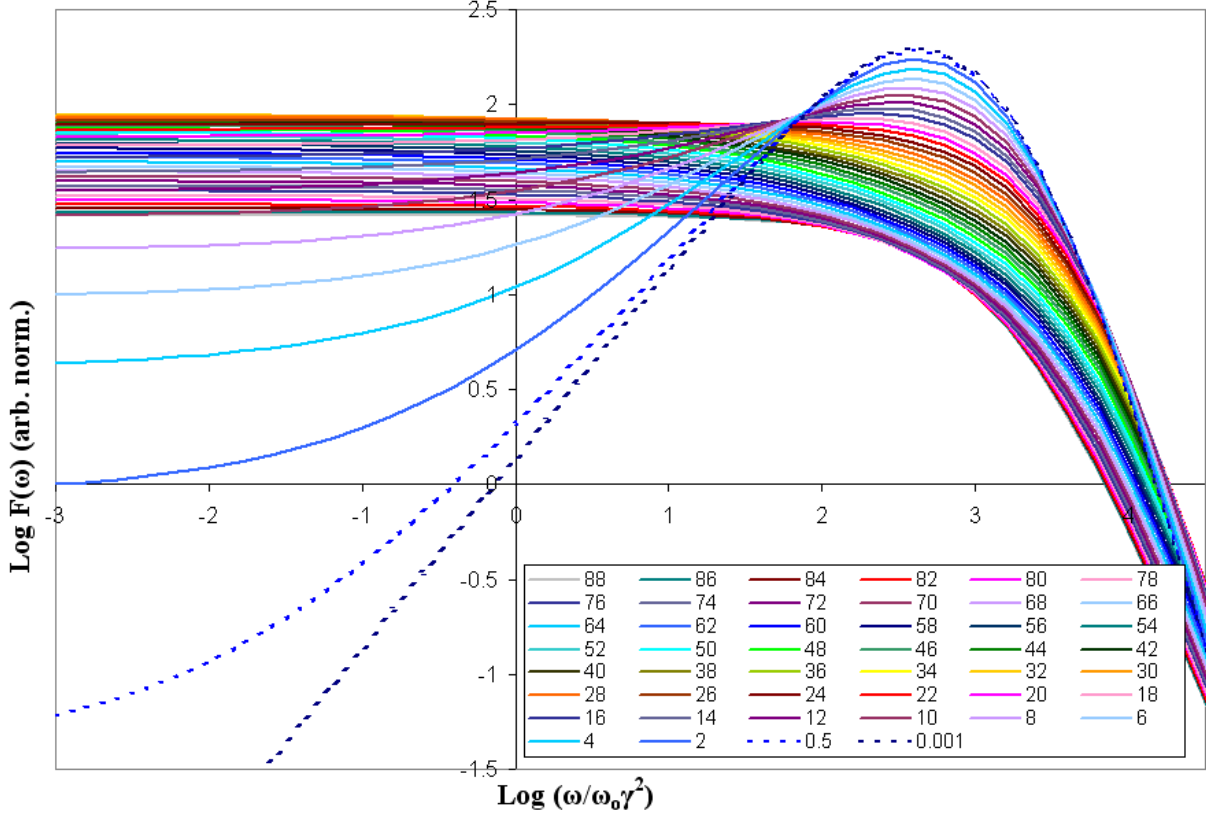


Fig. 9(a).— The angle-averaged radiation power spectra (F_ν) of jitter radiation, numerically calculated for a step size of 0.2 in $\log(\omega)$ for every two degrees in θ . The dotted lines are numerical calculations for values close to $\theta=0$, illustrating the behavior of the spectrum in this limit. The spectra are arbitrarily normalized so that the first calculated value for the $\theta = 2$ spectrum is 0. The radiation spectra are flat for low- ω , then slope upwards to a notable peak for low θ . As θ progresses to $\pi/2$, the low- ω' amplitude increases and the size of the sloped region decreases until the spectra becomes relatively flat until its sharp decline at high- ω .

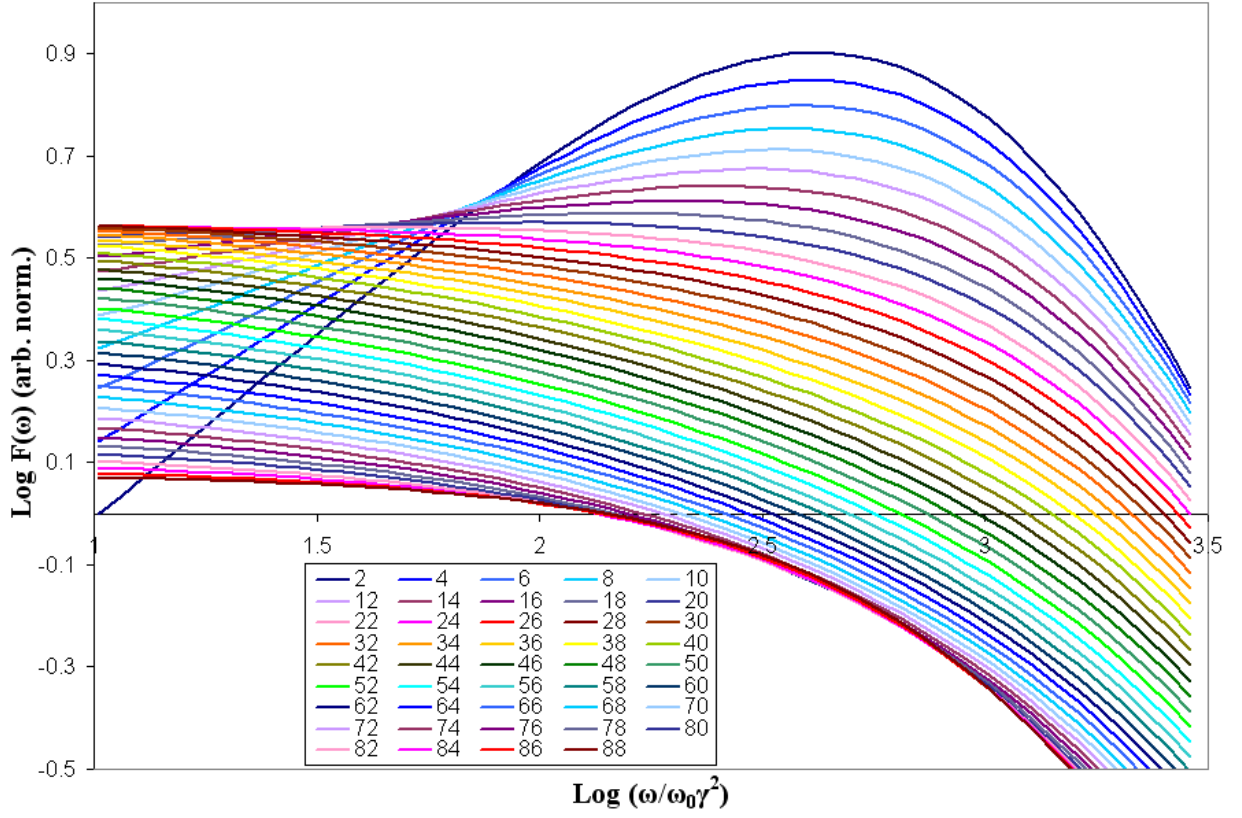


Fig. 9(b).— A more detailed view of the radiation spectra in the region of the spectral peak. The spectral values are calculated every 0.05 in $\log(\omega)$ for finer resolution of the detail in this region.

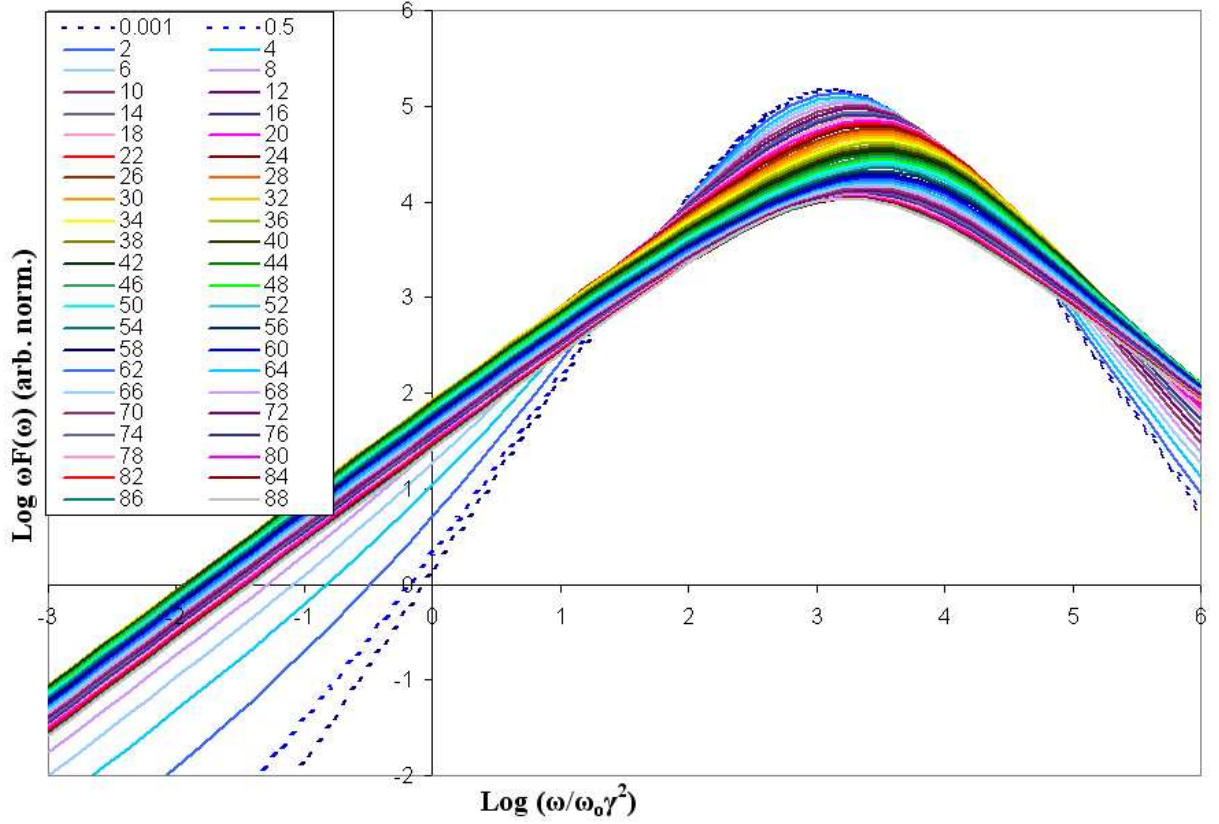


Fig. 10.— The νF_ν spectra calculated from our jitter radiation power spectral results. The νF_ν spectral peak is the peak energy E_p used in the Band functional fit commonly used for GRB spectra (Band et al. 1993; Kaneko et al. 2006).

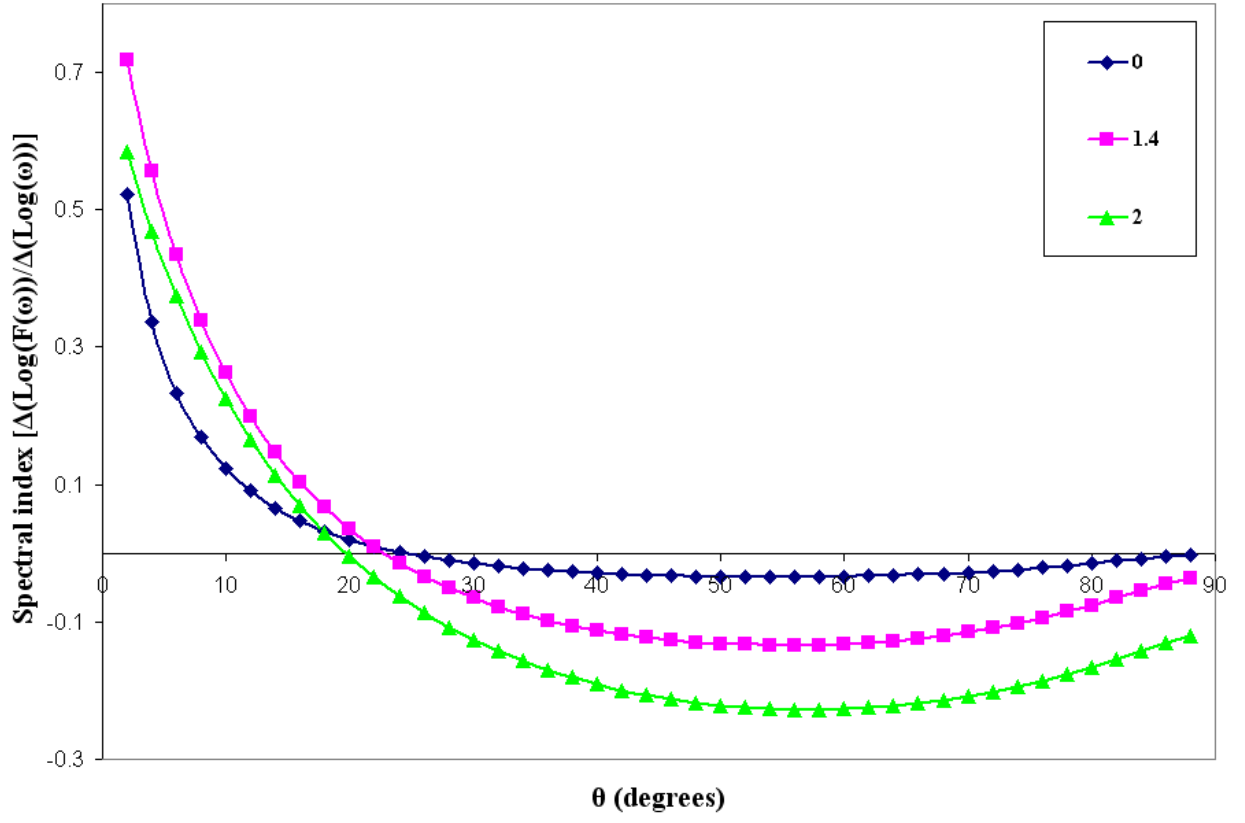


Fig. 11.— A plot of the slope of the radiation spectra taken at several values of $\log(\omega/\omega_0\gamma^2)$ over the range of θ .

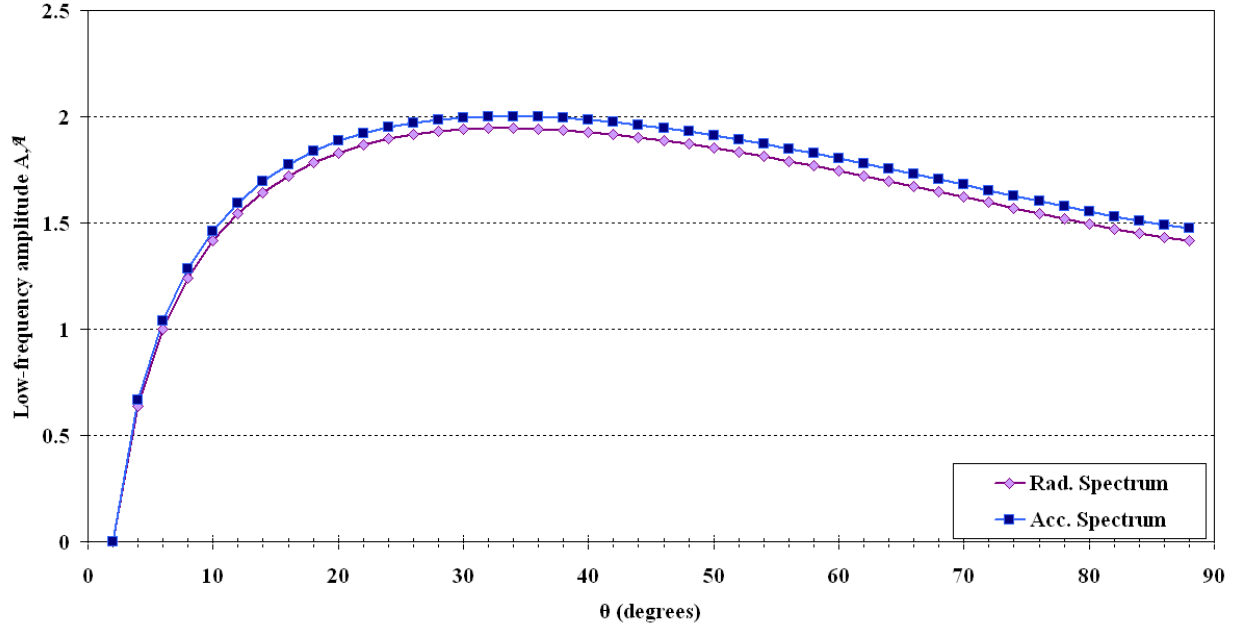


Fig. 12.— The angular dependence of the low-frequency amplitude A of our calculated jitter radiation spectra and the corresponding variation of the low-frequency amplitude \mathcal{A} of our calculated acceleration spectra. The low-energy amplitude in each case is taken to be the first calculated value of the spectrum; for our choice of parameters and calculation window, this initial value is well below the first spectral break for all $\theta > 2^\circ$. In both cases, we have normalized our spectra such that the low-energy amplitude at $\theta = 2^\circ$ is 0.

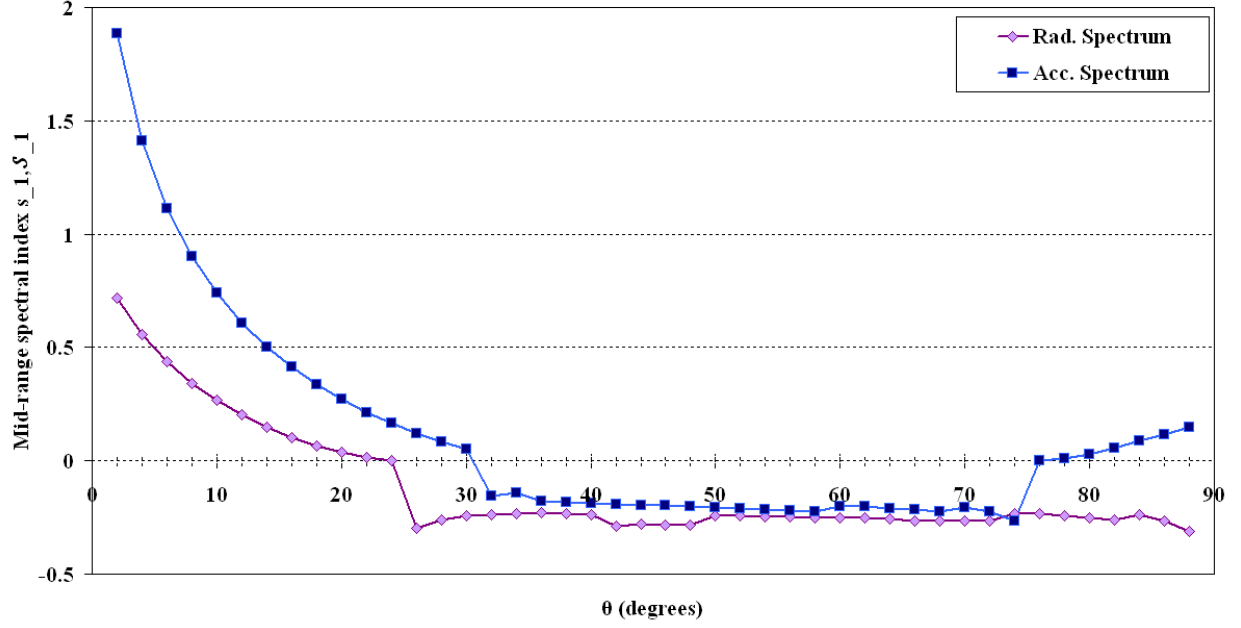


Fig. 13.— The angular dependence of the mid-range spectral indices s_1 of our calculated jitter radiation spectra and s_1 of our calculated acceleration spectra. For peaked spectra the mid-range spectral index is the maximum slope below the spectral peak; for unpeaked spectra the mid-range spectral index is the average slope between the point at which the spectrum falls below $A - 0.01$ and the point at which the numerical second derivative reaches its minimum value (the "drop point").

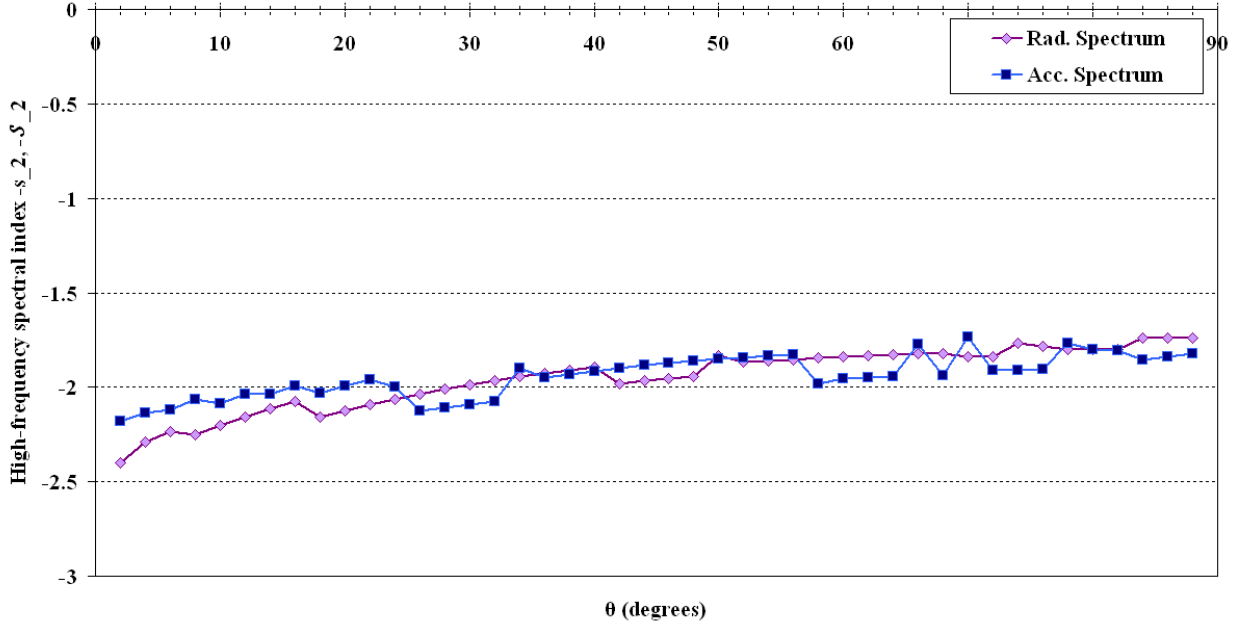


Fig. 14.— The angular dependence of the high-frequency spectral indices s_2 for our calculated jitter radiation spectra and s_2 of the corresponding acceleration spectra. The high-frequency spectral index is calculated as the slope between the drop point (the position of the largest negative change in slope) and the higher-frequency position at which the numerical second derivative is closest to 0.

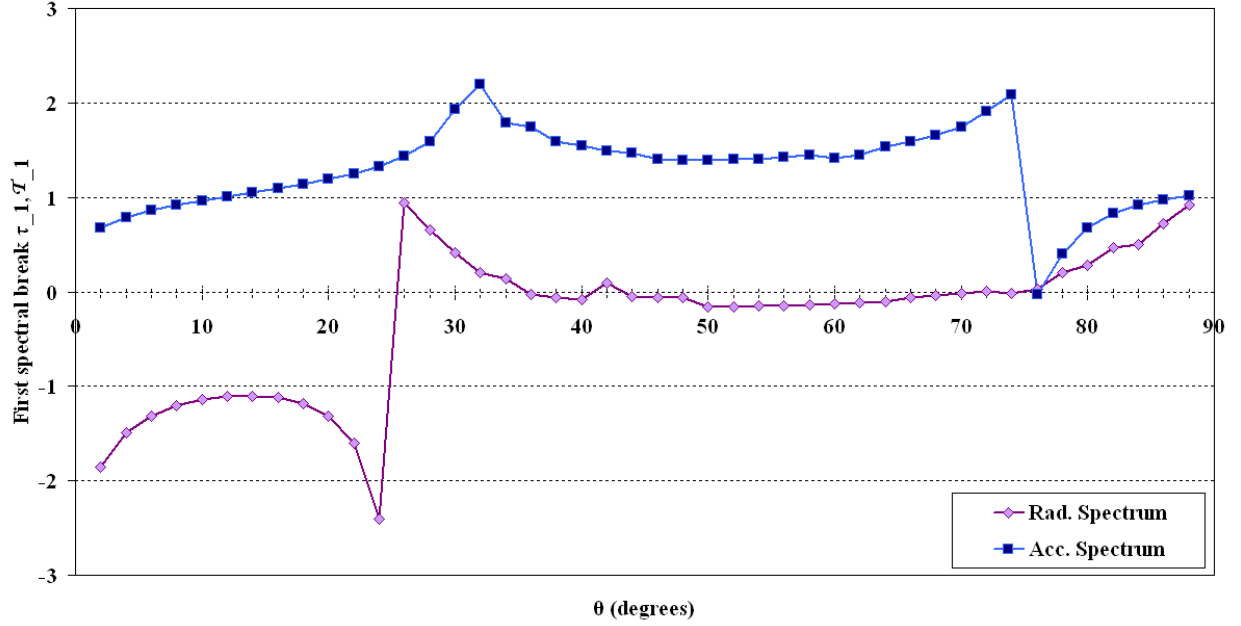


Fig. 15.— The angular dependence of the first spectral break positions $\log_e(\omega/\omega_o\gamma^2) = \tau_1$ and $\log_e(\omega'/\omega'_o) = \tau_1$ in our calculated jitter radiation spectra and acceleration spectra, respectively. In each case these transition points are found as the intersection between the low-frequency fit line of slope 0 and the mid-range fit line of slope s_1 (radiation) or s_1 (acceleration), as found by the fit described in detail in section 4.

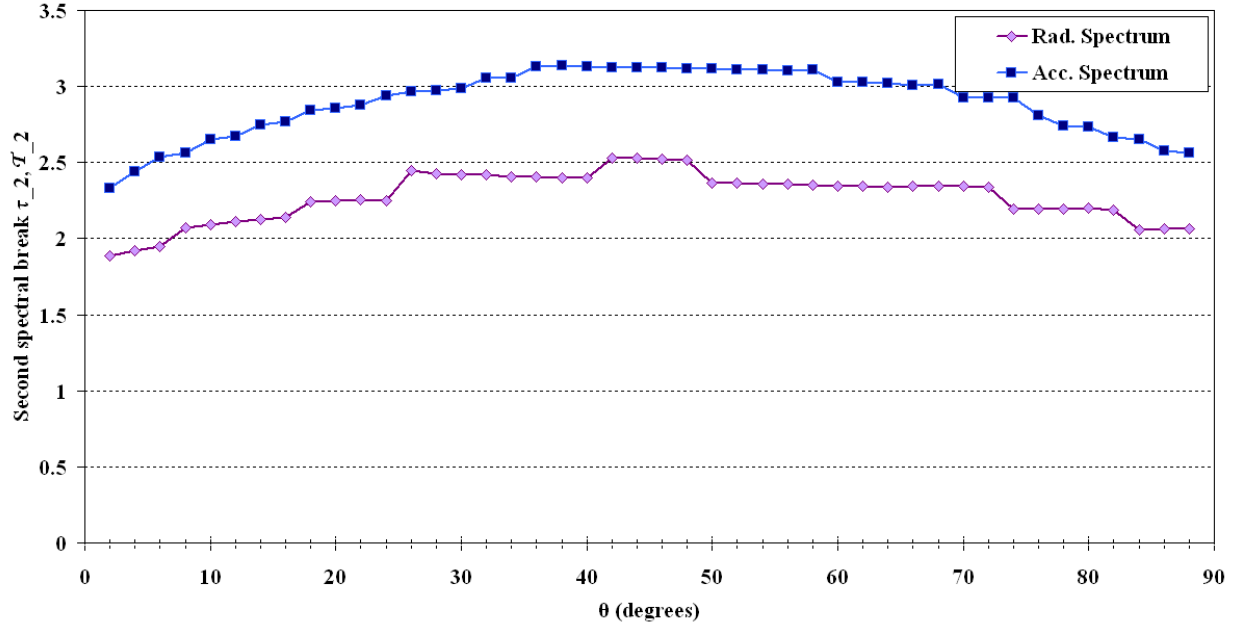


Fig. 16.— The angular dependence of the second spectral break τ_2 of the jitter radiation spectra and τ_2 of the acceleration spectra. The transition point is the calculated intersection between the mid-range fit line of slope s_1 (radiation) or s_1 (acceleration) and the high-frequency fit line of slope $-s_2$ (radiation) or $-s_2$ (acceleration).

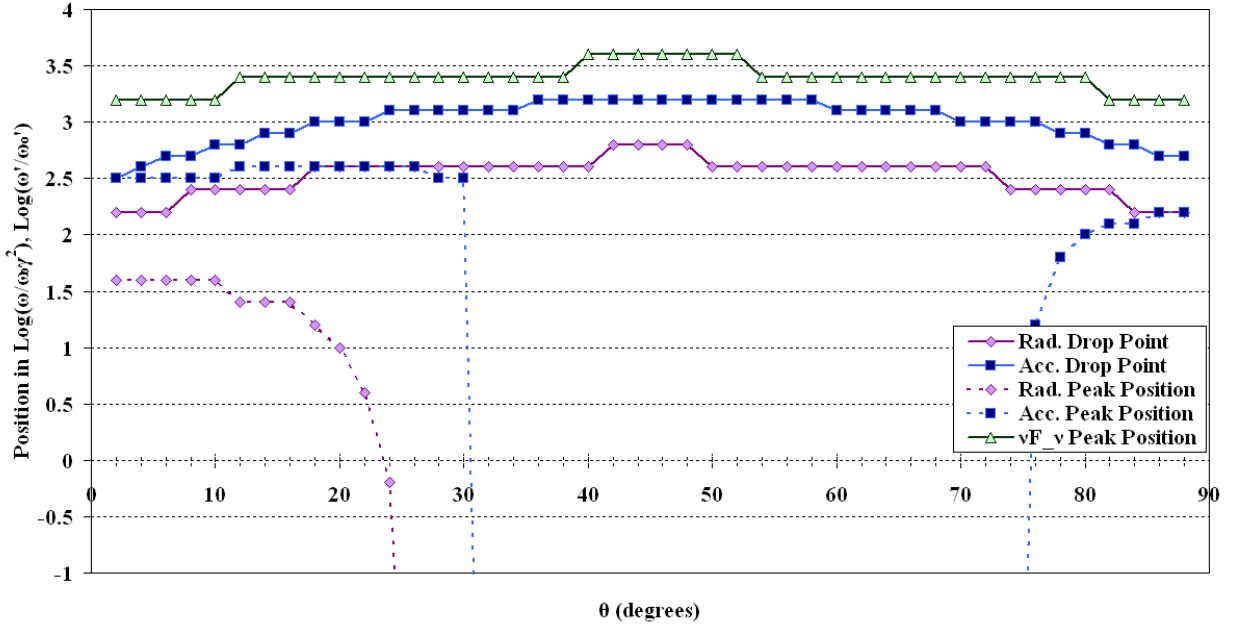


Fig. 17.— A comparison of the angular dependence of the positions of the radiation and acceleration spectral peaks and the spectral drop points, where the numerical second derivatives of our calculated spectra reach a minimum. We have also plotted the angular dependence of the peak in the νF_ν spectrum. The drop point in our F_ν radiation spectrum nicely tracks the behavior of the peak in the νF_ν spectrum, which is peak energy E_p in the Band function (Band et al. 1993) commonly used to fit GRB spectra.

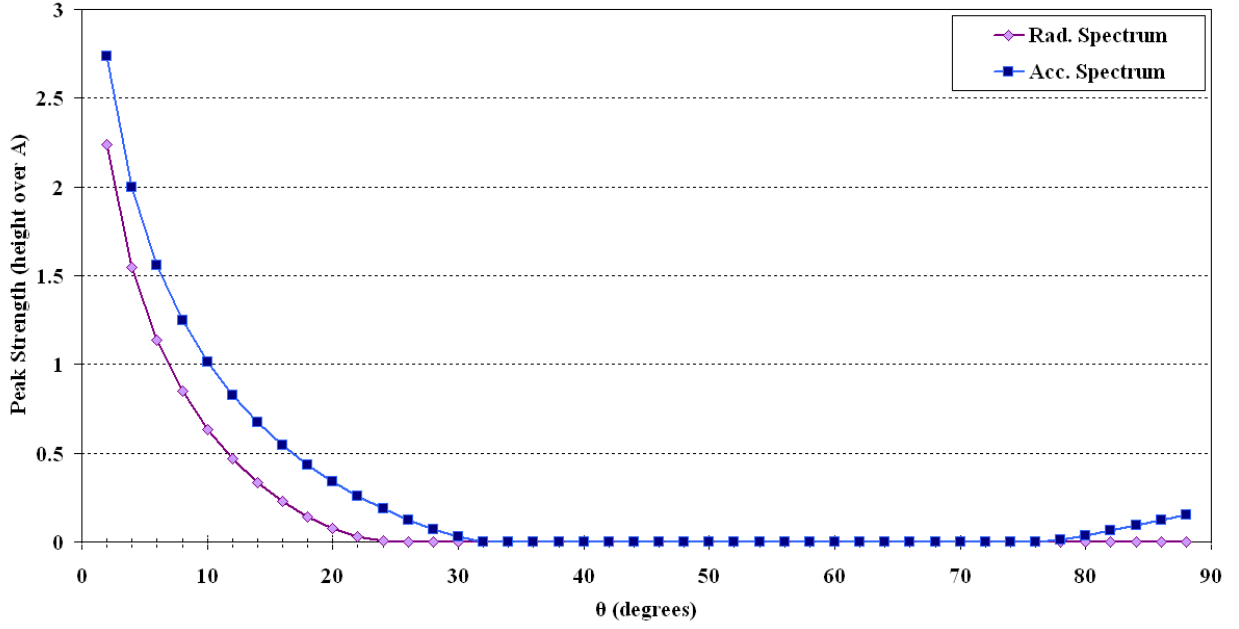


Fig. 18.— The angular dependence of the strength of the spectral peak, i.e. the height of the peak above the initial low-frequency spectral amplitude A . We see that the peak disappears in the radiation spectrum at θ roughly 8° less than in the acceleration spectrum, and does not reappear at θ close to 90° .

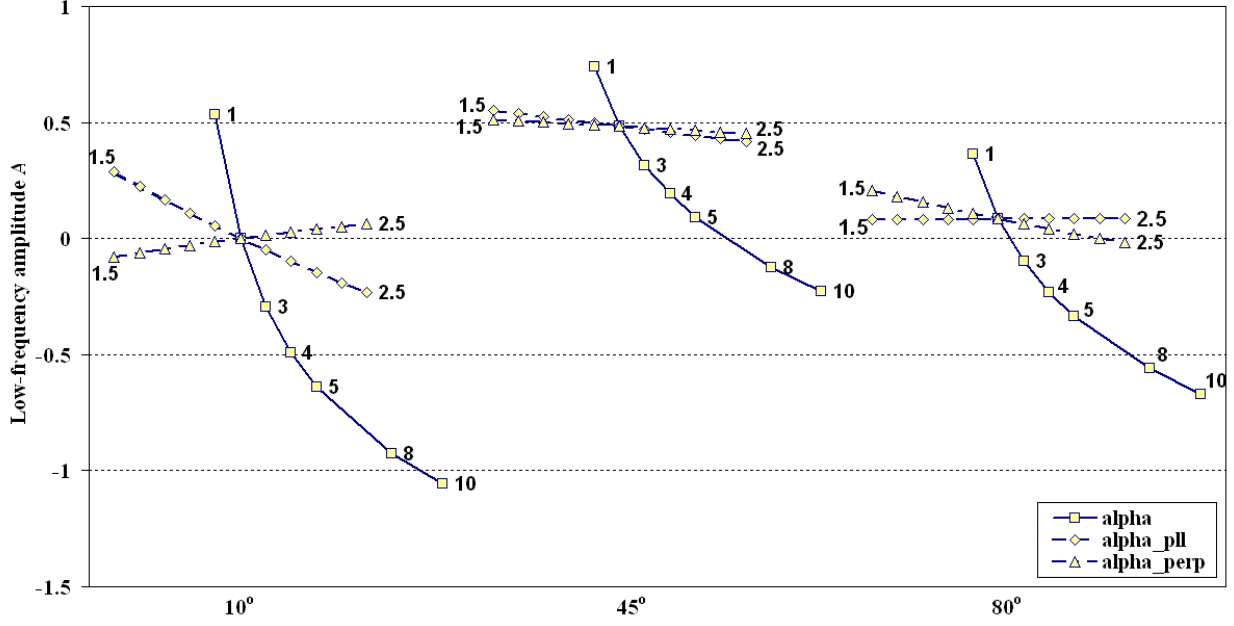


Fig. 19(a).— This figure shows the variation of the low-frequency amplitude of the radiation spectrum (vertical axis) with changes in the magnetic field parameters α_i , which are varied progressively in the horizontal dimension as described below, for each of the fixed viewing angles $\theta = 10^\circ$, 45° , and 80° (first, second, and third clusters respectively). The solid line with square data points indicates the behavior when α is jointly varied (over a range from 1 to 10, as indicated) in both the transverse and parallel magnetic field equation: $\alpha = \alpha_\perp = \alpha_\parallel$. The dotted lines show the effect of individually varying α_\perp (triangular data points) and α_\parallel (diamond data points), which characterize the magnetic field transverse to and along the instability filamentation axis, respectively. The parameters α_\perp and α_\parallel are varied from a starting value of 1.5 to an ending value of 2.5, in increments of 0.1. Our original value of $\alpha = \alpha_\perp = \alpha_\parallel = 2.0$ is the central data point where the lines intersect.

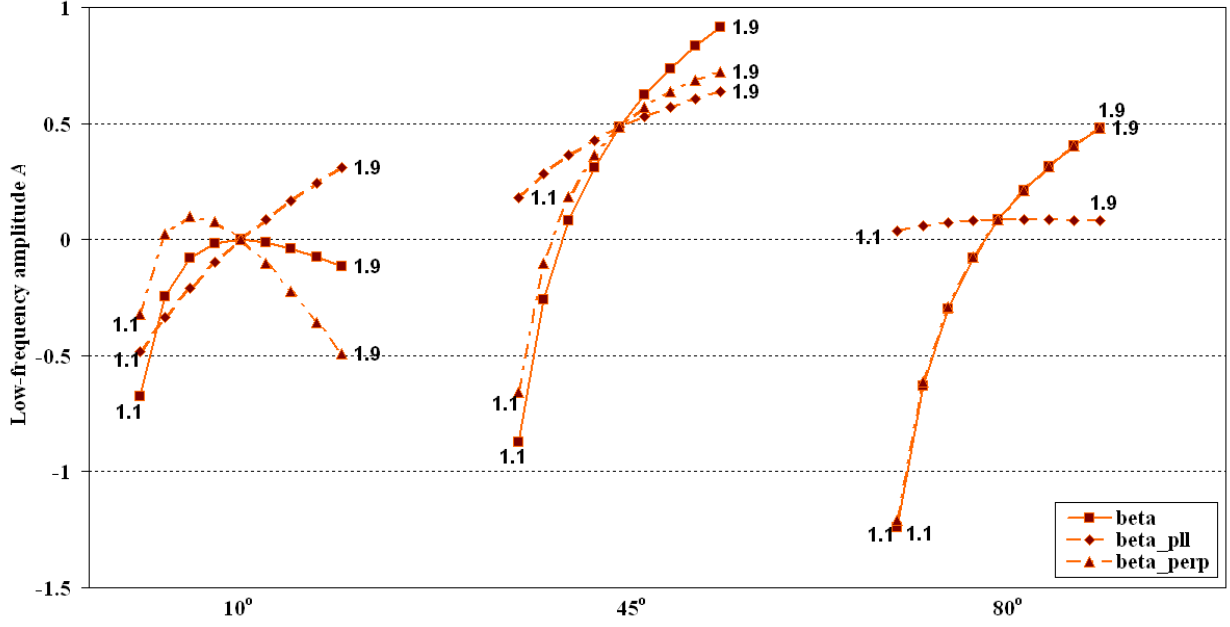


Fig. 19(b).— This figure shows the variation of the low-frequency amplitude of the radiation spectrum (vertical axis) with changes in the magnetic field parameters β_i , which are varied progressively in the horizontal dimension for the fixed viewing angles $\theta = 10^\circ$, 45° , and 80° (first, second, and third clusters respectively). The solid line with square data points indicates the behavior when β is jointly varied from 1.1 to 1.9, in increments of 0.1, in both the transverse and parallel magnetic field equation: $\beta = \beta_\perp = \beta_\parallel$. The dotted lines show the effect of varying β_\perp (triangular data points) and β_\parallel (diamond data points) individually, also from 1.1 to 1.9, in increments of 0.1. Our original value of $\beta = \beta_\perp = \beta_\parallel = 1.5$ is the central data point where the lines intersect.

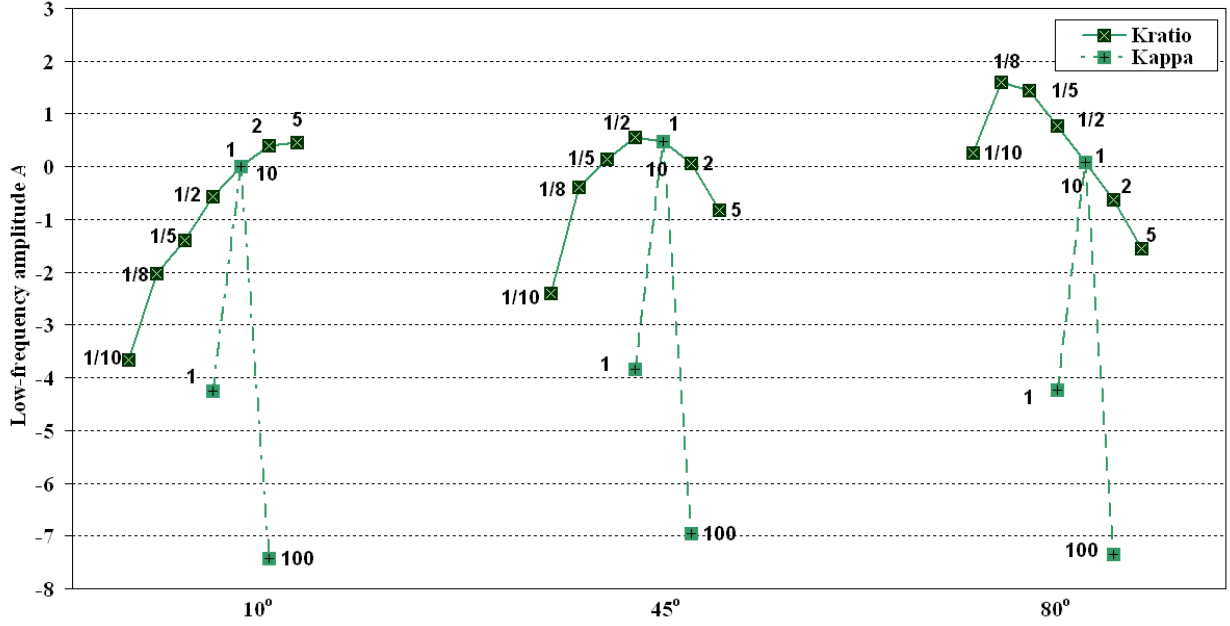


Fig. 19(c).— This figure shows the variation of the low-frequency amplitude of the radiation spectrum with changes in the magnetic field parameters κ , which are varied progressively in the horizontal dimension for the fixed viewing angles $\theta = 10^\circ$, 45° , and 80° (first, second, and third clusters respectively). We vary κ jointly ($\kappa = \kappa_\perp = \kappa_\parallel$) by powers of 10, from 1 to 100, as indicated. We also vary κ_\perp and κ_\parallel relative to one another by changing the ratio $K = \kappa_\perp/\kappa_\parallel$ through a range of values as indicated.

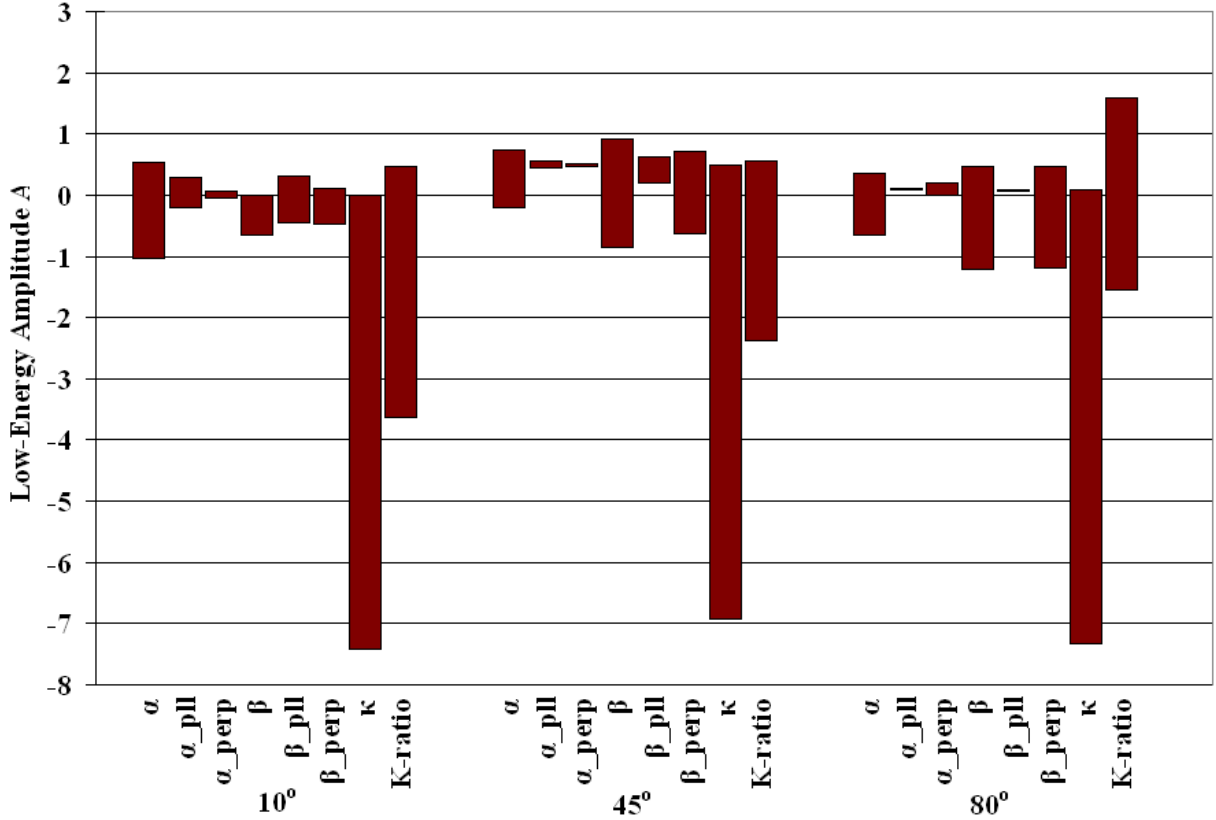


Fig. 20.— This figure compares the influences of the magnetic field spectral parameter variations on the low-frequency amplitude A of the radiation spectrum obtained for representative viewing angles $\theta = 10^\circ$, 45° , and 80° . For each spectral parameter (indicated on the bottom axis) the graph indicates the range between the maximum and minimum values of A obtained by our variations of that parameter. (The parameter variations are as indicated in the previous figures and described in detail in section 5.)

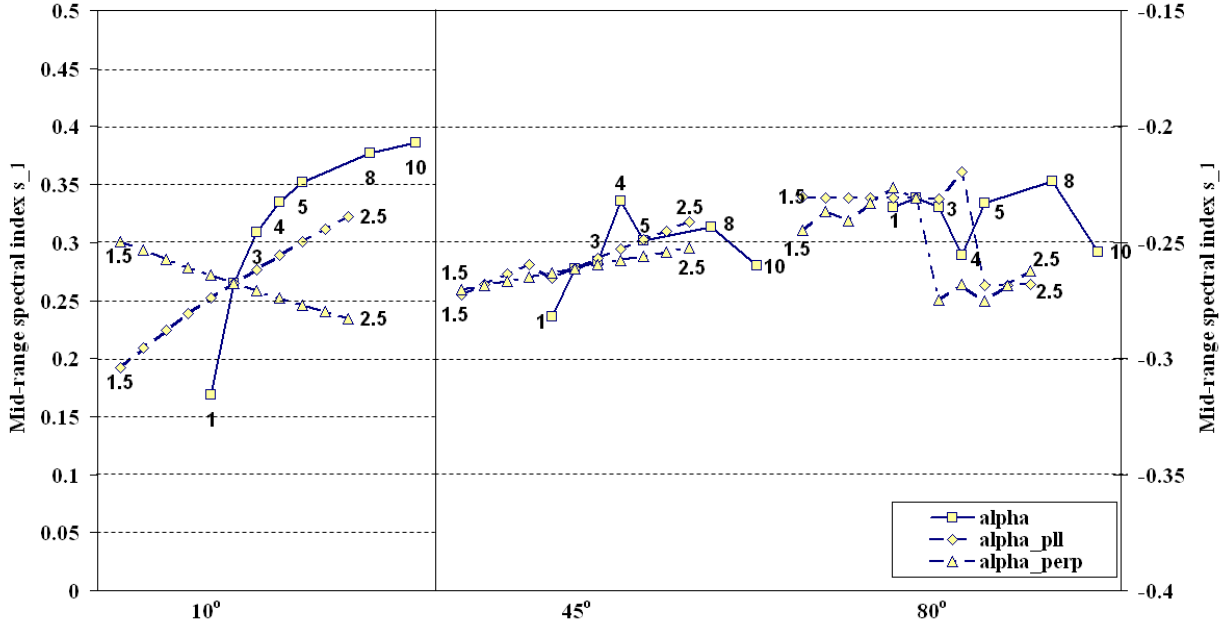


Fig. 21(a).— This figure shows the variation of the mid-range spectral index (as determined by the fit described in section 5) with changes in the magnetic field parameters α_i . This is the maximum slope below the peak for peaked spectra, and the average slope in intermediary region for unpeaked spectra, and is shown for representative viewing angles $\theta = 10^\circ$, 45° , and 80° . The solid line with square data points indicates the behavior when α is jointly varied (over a range from 1 to 10, as indicated) in both the transverse and parallel magnetic field equation: $\alpha = \alpha_{\perp} = \alpha_{||}$. The dotted lines show the effect of varying α_{\perp} (triangular data points) and $\alpha_{||}$ (diamond data points) individually, from 1.5 to 2.5, in increments of 0.1.

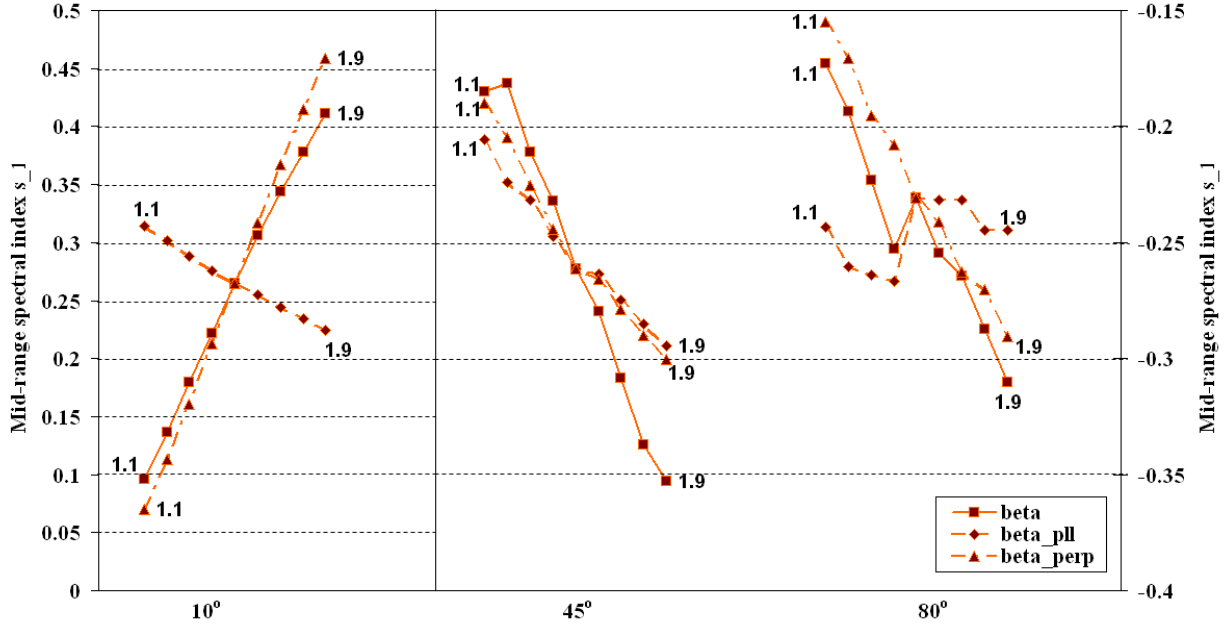


Fig. 21(b).— This figure shows the variation of the mid-range spectral index with changes in the magnetic field parameters β_i . The solid line with square data points indicates the behavior when β is jointly varied from 1.1 to 1.9, in increments of 0.1, in both the transverse and parallel magnetic field equation: $\beta = \beta_{\perp} = \beta_{||}$. The dotted lines show the effect of varying β_{\perp} (triangular data points) and $\beta_{||}$ (diamond data points) individually, also from 1.1 to 1.9, in increments of 0.1.

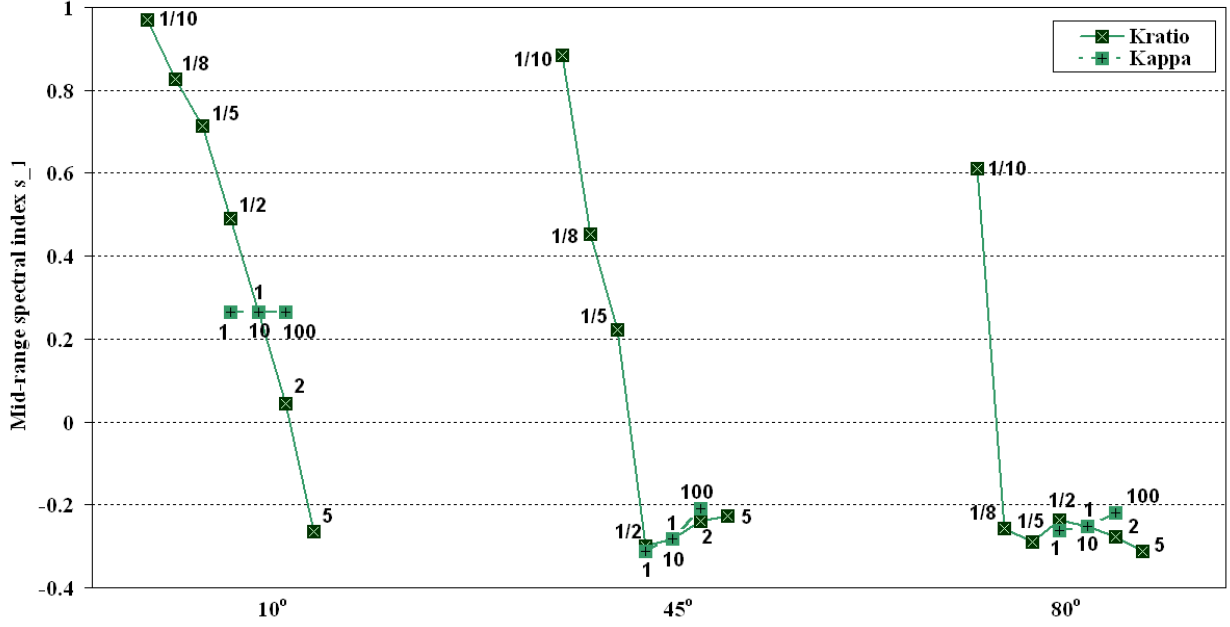


Fig. 21(c).— This figure shows the variation of the mid-range spectral index with changes in the magnetic field parameters κ . We vary κ jointly ($\kappa = \kappa_{\perp} = \kappa_{\parallel}$) by powers of 10, from 1 to 100, as indicated by the dotted line. We also vary κ_{\perp} and κ_{\parallel} relative to one another by changing the ratio $K = \kappa_{\perp}/\kappa_{\parallel}$ through a range of values as indicated by the solid line.

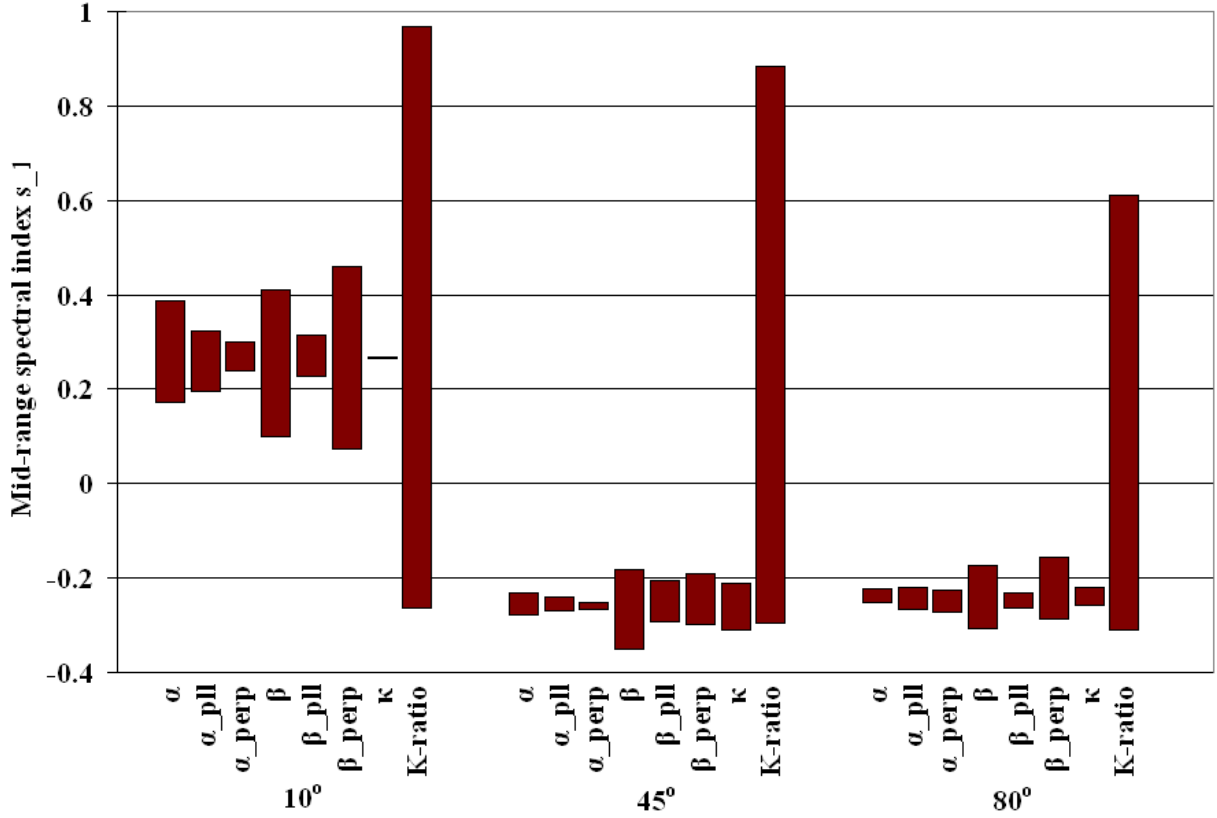


Fig. 22.— This figure compares the influences of different magnetic field spectral parameters on the mid-range spectral index s_1 of the radiation spectrum obtained for representative viewing angles $\theta = 10^\circ$, 45° , and 80° . For each spectral parameter (indicated on the bottom axis) the graph indicates the range between the maximum and minimum values of s_1 obtained by our variations of that parameter (The parameter variations are as indicated in the previous figure and described in detail in section 5.)

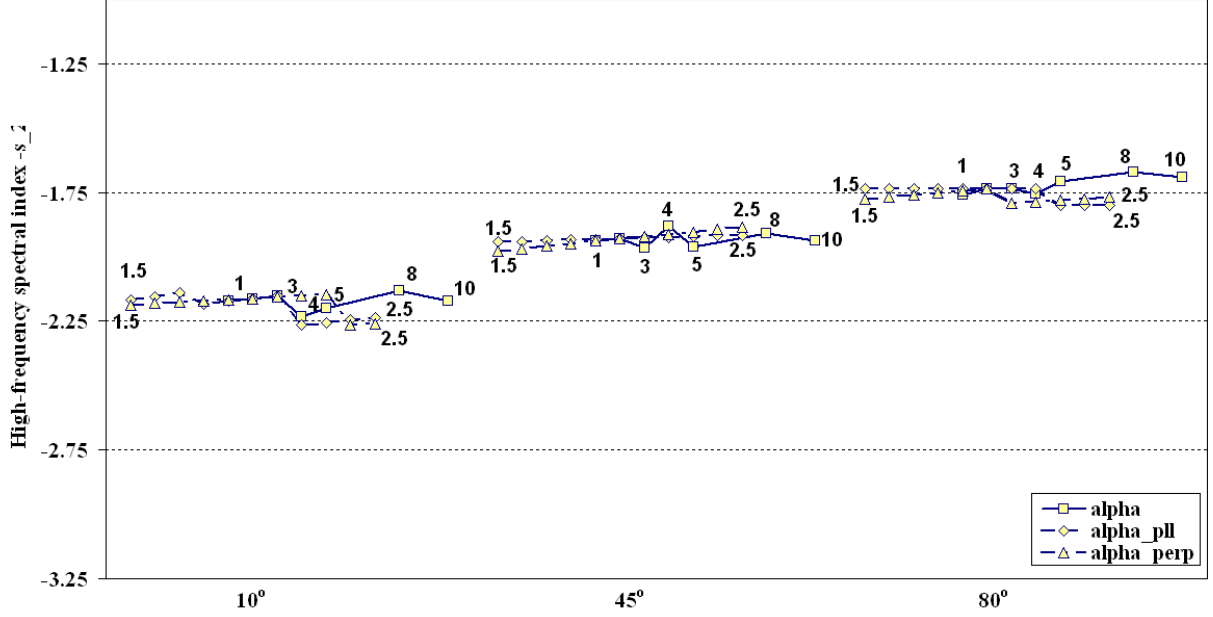


Fig. 23(a).— This figure shows the variation of the high-frequency spectral index (as determined by the fit described in section 5) with changes in the magnetic field parameters α_i . The solid line with square data points indicates the behavior when α is jointly varied (over a range from 1 to 10, as indicated) in both the transverse and parallel magnetic field equation: $\alpha = \alpha_{\perp} = \alpha_{\parallel}$. The dotted lines show the effect of varying α_{\perp} (triangular data points) and α_{\parallel} (diamond data points) individually, from 1.5 to 2.5, in increments of 0.1.

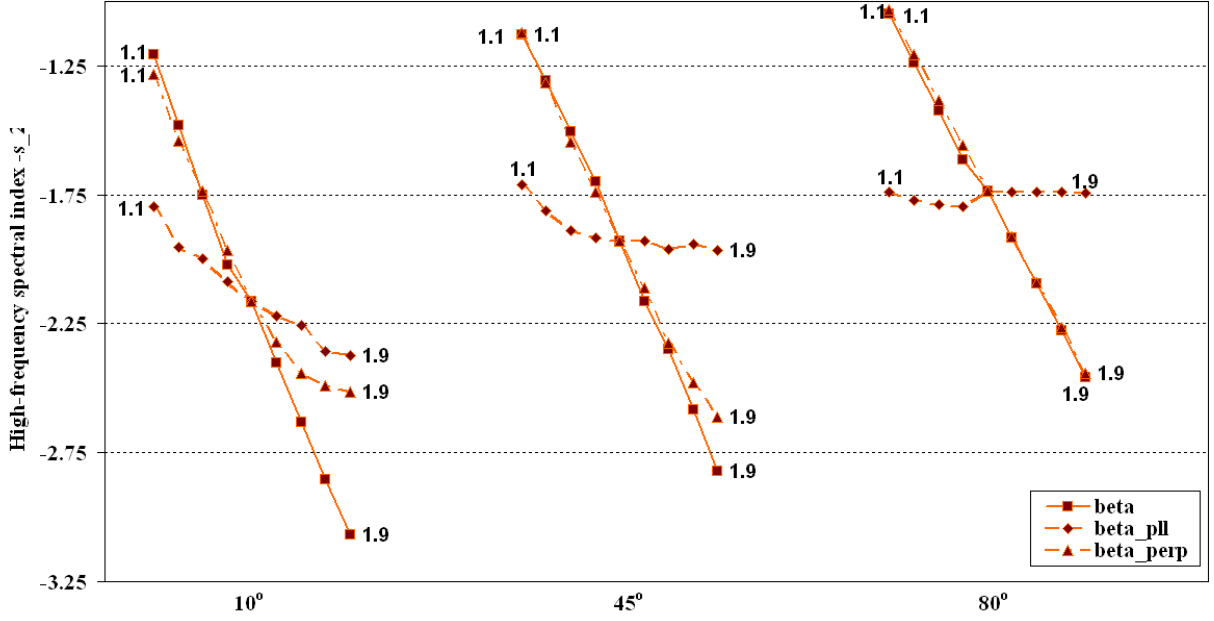


Fig. 23(b).— This figure shows the variation of the high-frequency spectral index with changes in the magnetic field parameters β_i . The solid line with square data points indicates the behavior when β is jointly varied from 1.1 to 1.9, in increments of 0.1, in both the transverse and parallel magnetic field equation: $\beta = \beta_{\perp} = \beta_{\parallel}$. The dotted lines show the effect of varying β_{\perp} (triangular data points) and β_{\parallel} (diamond data points) individually, also from 1.1 to 1.9, in increments of 0.1.

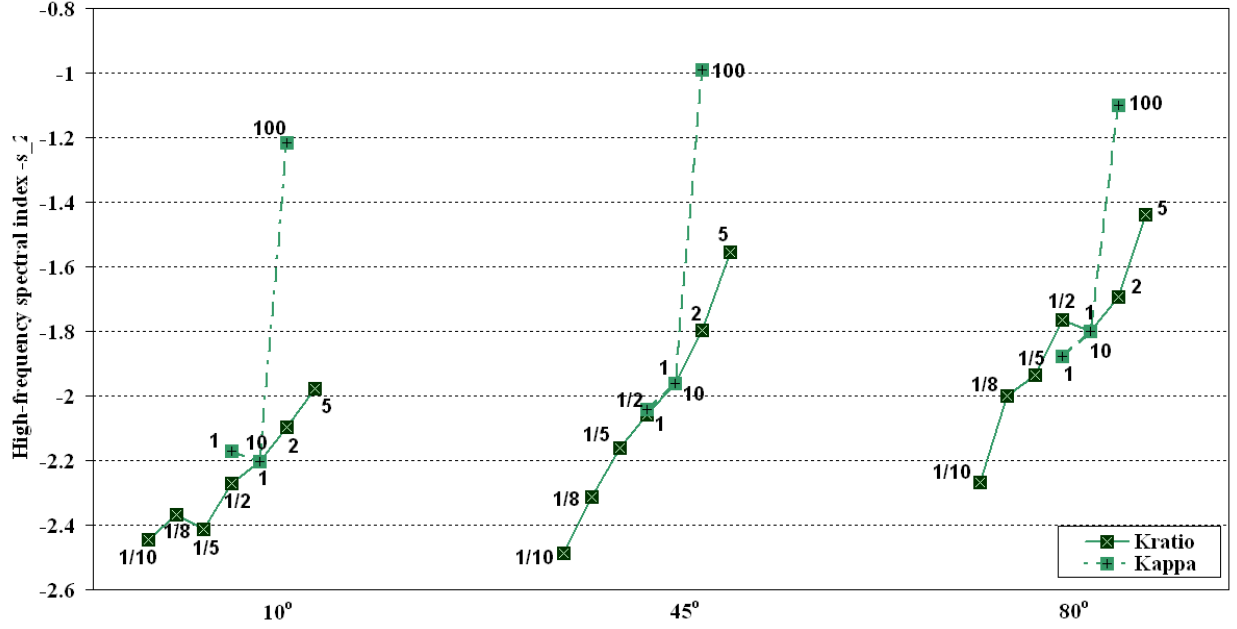


Fig. 23(c).— This figure shows the variation of the high-frequency spectral index with changes in the magnetic field parameters κ . We vary κ jointly ($\kappa = \kappa_{\perp} = \kappa_{\parallel}$) by powers of 10, from 1 to 100, as indicated. We also vary κ_{\perp} and κ_{\parallel} relative to one another by changing the ratio $K = \kappa_{\perp}/\kappa_{\parallel}$ through a range of values as indicated.

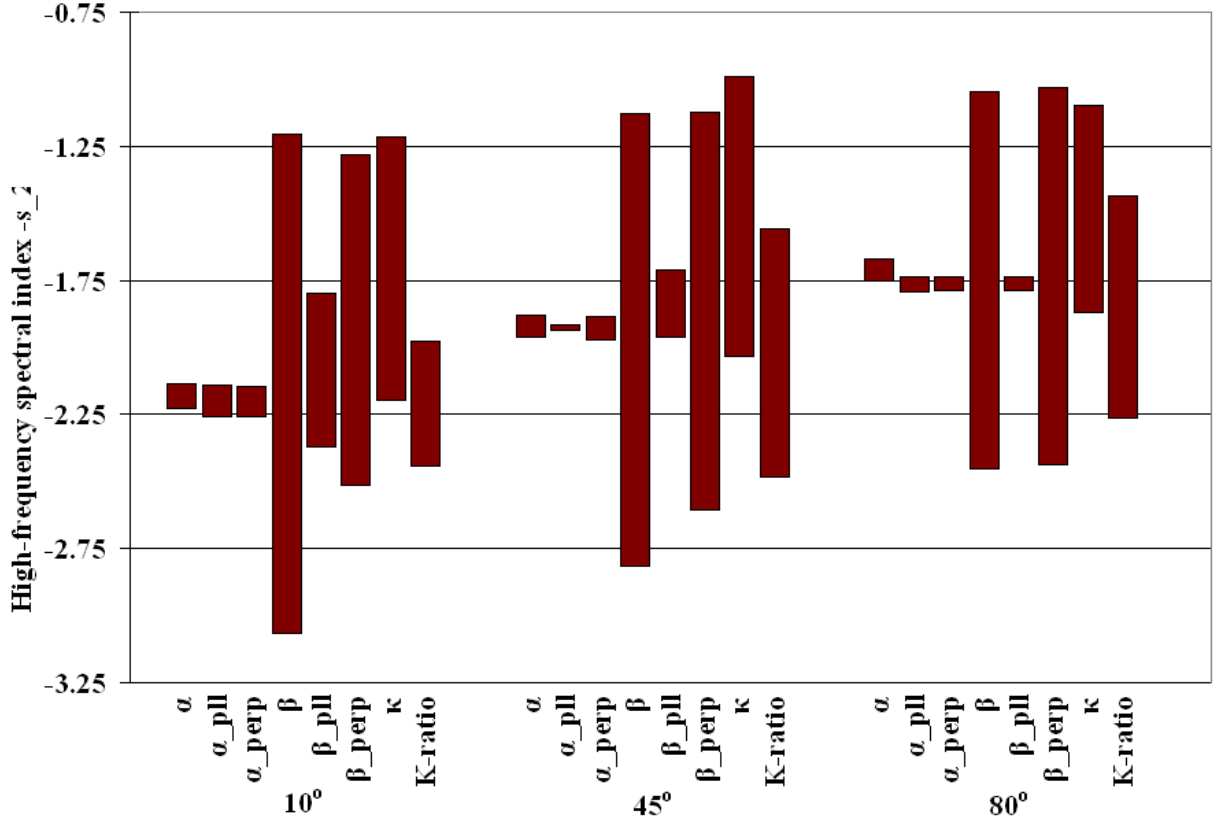


Fig. 24.— This figure compares the influences of different magnetic field spectral parameters on the high-frequency spectral index $-s_2$ of the radiation spectrum. For each spectral parameter (indicated on the bottom axis) the graph indicates the range between the maximum and minimum values of $-s_2$ obtained by our variations of that parameter (The parameter variations are as indicated in the previous figures and described in detail in section 5.)

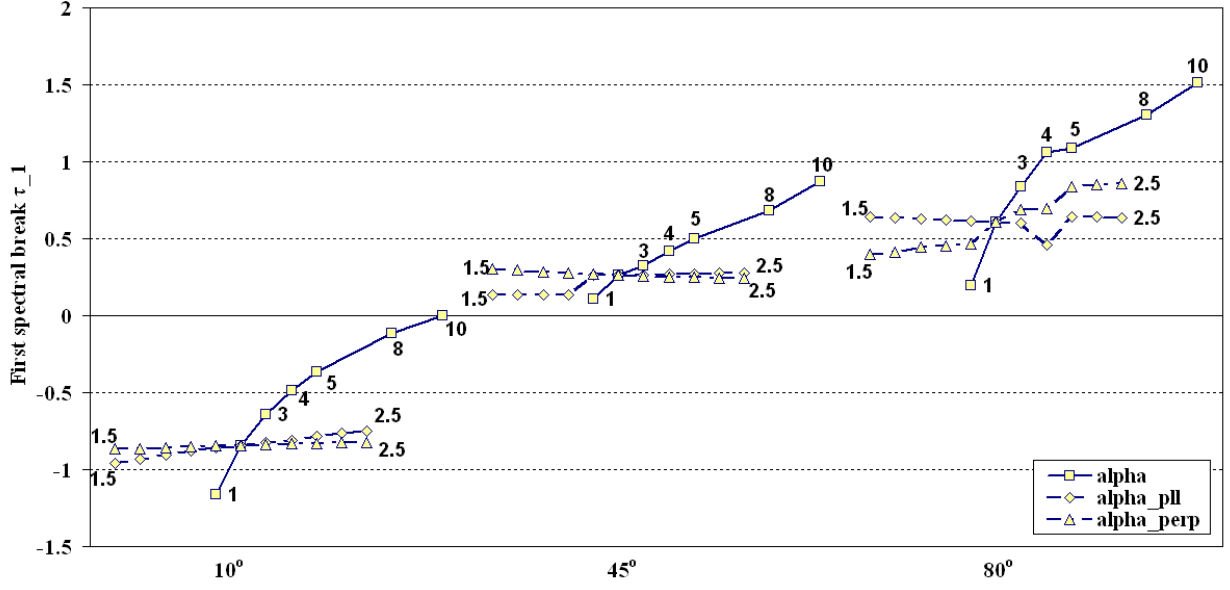


Fig. 25(a).— This figure shows the variation of the first spectral transition point (as described in section 5) with changes in the magnetic field parameters α_i , for $\theta = 10^\circ$, 45° , and 80° . The solid line with square data points indicates the behavior when α is jointly varied (over a range from 1 to 10, as indicated) in both the transverse and parallel magnetic field equation: $\alpha = \alpha_\perp = \alpha_\parallel$. The dotted lines show the effect of varying α_\perp (triangular data points) and α_\parallel (diamond data points) individually, from 1.5 to 2.5, in increments of 0.1.

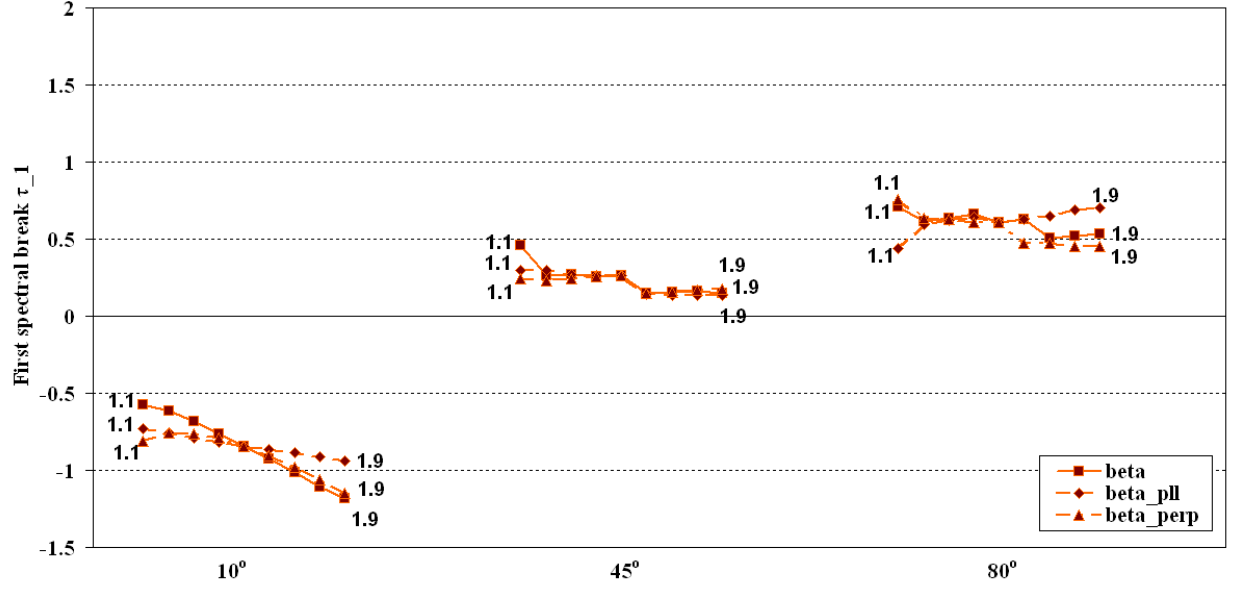


Fig. 25(b).— This figure shows the variation of the first spectral transition point with changes in the magnetic field parameters β_i , for $\theta = 10^\circ$, 45° , and 80° . The solid line with square data points indicates the behavior when β is jointly varied from 1.1 to 1.9, in increments of 0.1, in both the transverse and parallel magnetic field equation: $\beta = \beta_\perp = \beta_\parallel$. The dotted lines show the effect of varying β_\perp (triangular data points) and β_\parallel (diamond data points) individually, also from 1.1 to 1.9, in increments of 0.1.

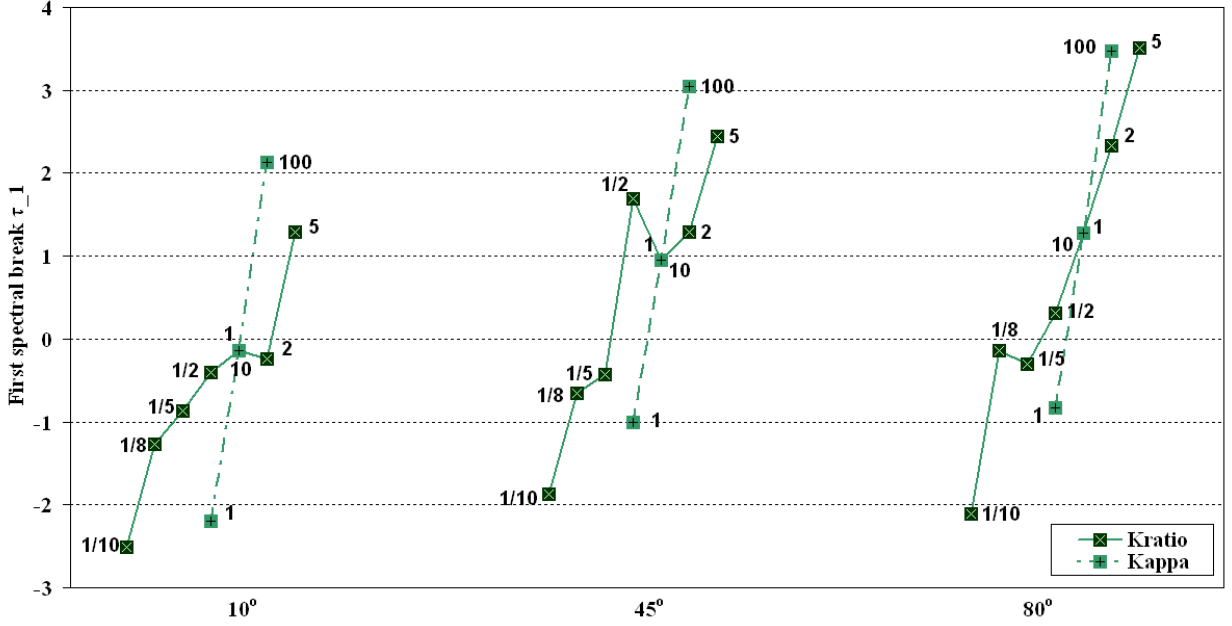


Fig. 25(c).— This figure shows the variation of the first spectral transition point with changes in the magnetic field parameters κ_i , for $\theta = 10^\circ$, 45° , and 80° . We vary κ jointly ($\kappa = \kappa_\perp = \kappa_\parallel$) by powers of 10, from 1 to 100, as indicated by the dotted lines. We also vary κ_\perp and κ_\parallel relative to one another by changing the ratio $K = \kappa_\perp/\kappa_\parallel$ through a range of values as indicated by the solid lines.

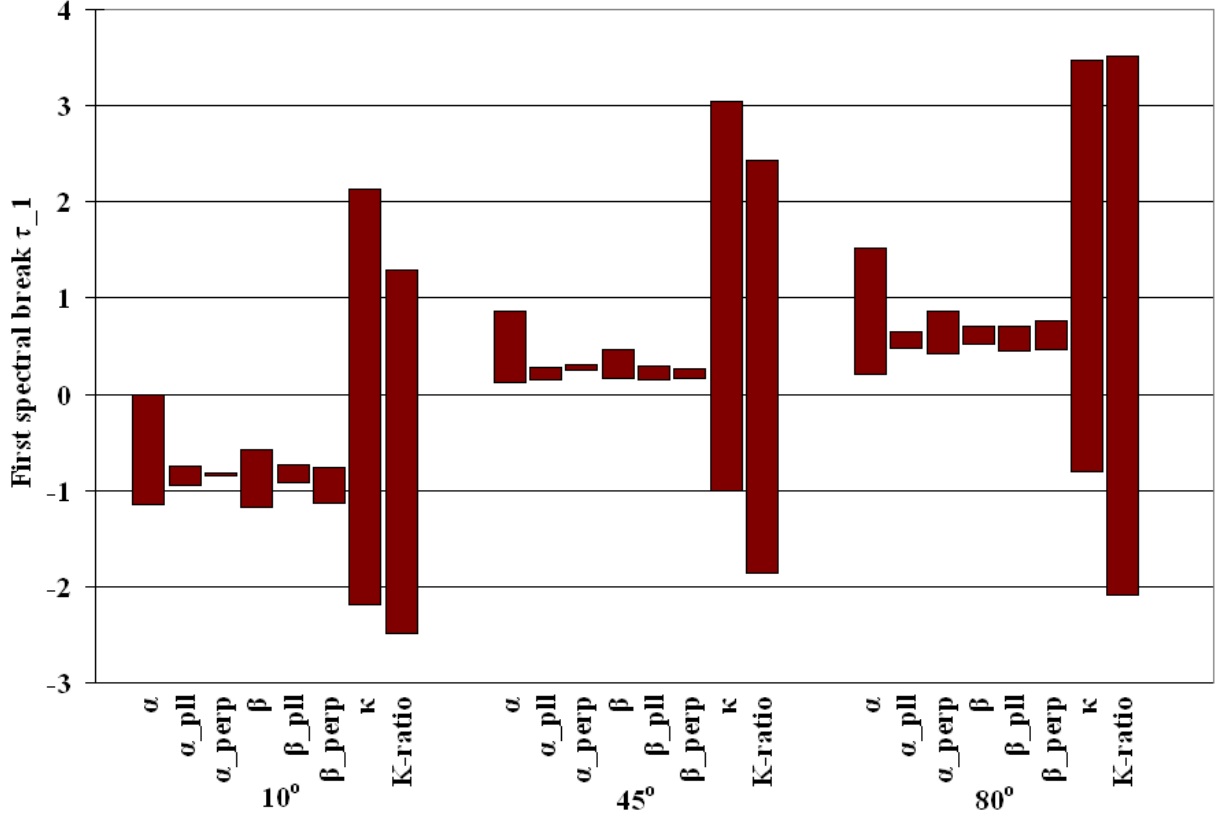


Fig. 26.— This figure compares the influences of different magnetic field spectral parameters on the first spectral break $\log_e(\omega/\omega_o\gamma^2) = \tau_1$ of the radiation spectrum obtained for representative viewing angles $\theta = 10^\circ$, 45° , and 80° . For each spectral parameter (indicated on the bottom axis) the graph indicates the range between the maximum and minimum values of τ_1 obtained by our variations of that parameter. (The parameter variations are as indicated in the previous figure and described in detail in section 5.)

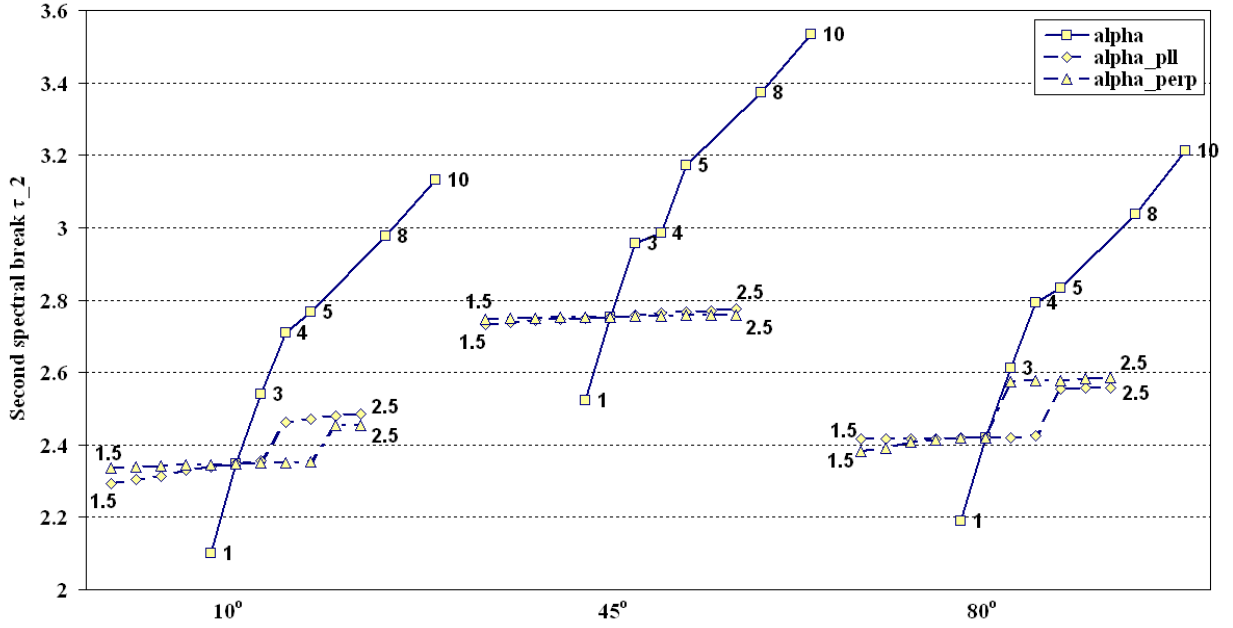


Fig. 27(a).— This figure shows the variation of the second spectral break (as determined by the fit described in section 5) with changes in the magnetic field parameters α_i . The solid line with square data points indicates the behavior when α is jointly varied (over a range from 1 to 10, as indicated) in both the transverse and parallel magnetic field equation: $\alpha = \alpha_{\perp} = \alpha_{\parallel}$. The dotted lines show the effect of varying α_{\perp} (triangular data points) and α_{\parallel} (diamond data points) individually, from 1.5 to 2.5, in increments of 0.1.

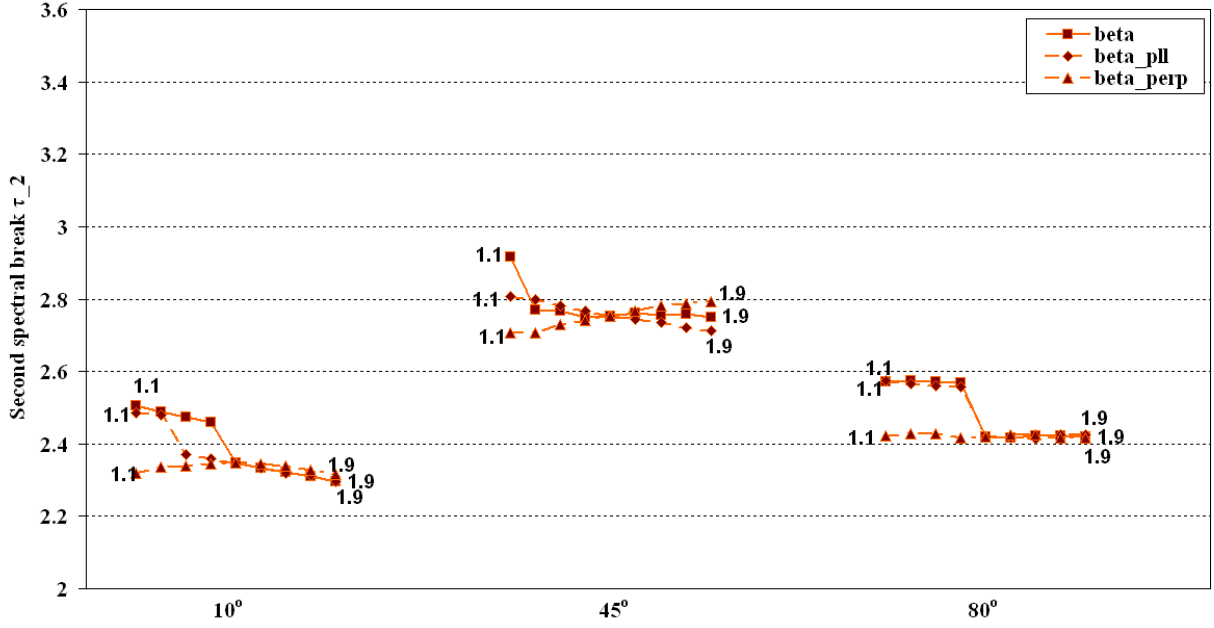


Fig. 27(b).— This figure shows the variation of the second spectral break with changes in the magnetic field parameters β_i . The solid line with square data points indicates the behavior when β is jointly varied from 1.1 to 1.9, in increments of 0.1, in both the transverse and parallel magnetic field equation: $\beta = \beta_{\perp} = \beta_{\parallel}$. The dotted lines show the effect of varying β_{\perp} (triangular data points) and β_{\parallel} (diamond data points) individually, also from 1.1 to 1.9, in increments of 0.1.

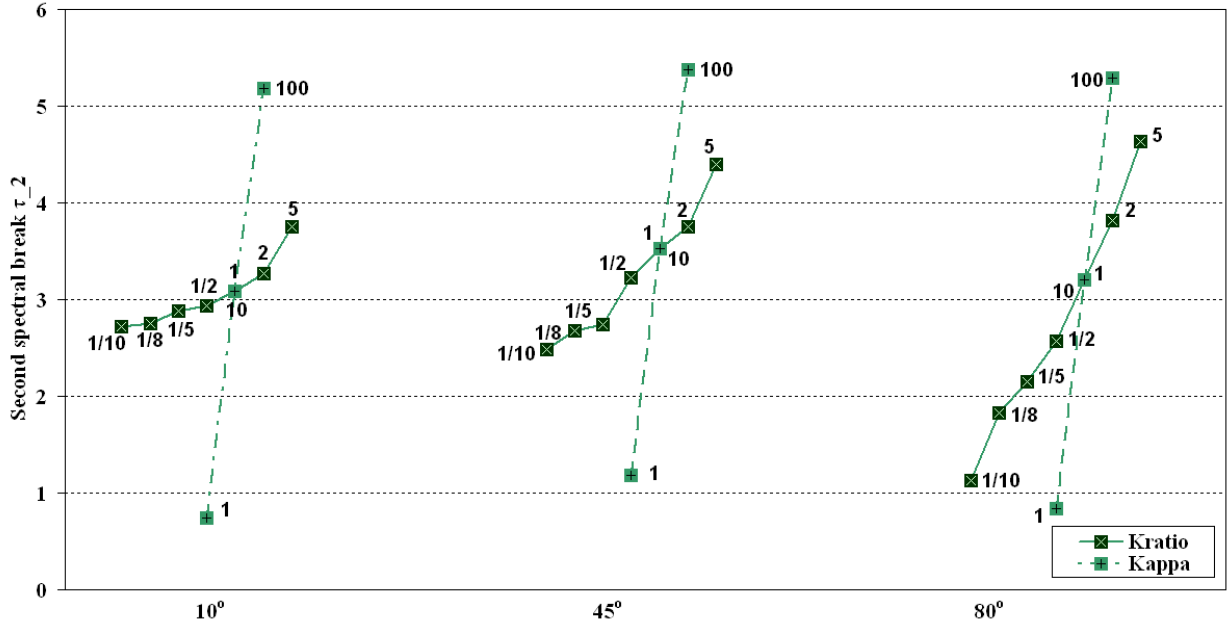


Fig. 27(c).— This figure shows the variation of the second spectral break with changes in the magnetic field parameters κ . We vary κ jointly ($\kappa = \kappa_\perp = \kappa_\parallel$) by powers of 10, from 1 to 100, as indicated by the dotted line. We also vary κ_\perp and κ_\parallel relative to one another by changing the ratio $K = \kappa_\perp/\kappa_\parallel$ through a range of values as indicated by the solid line.

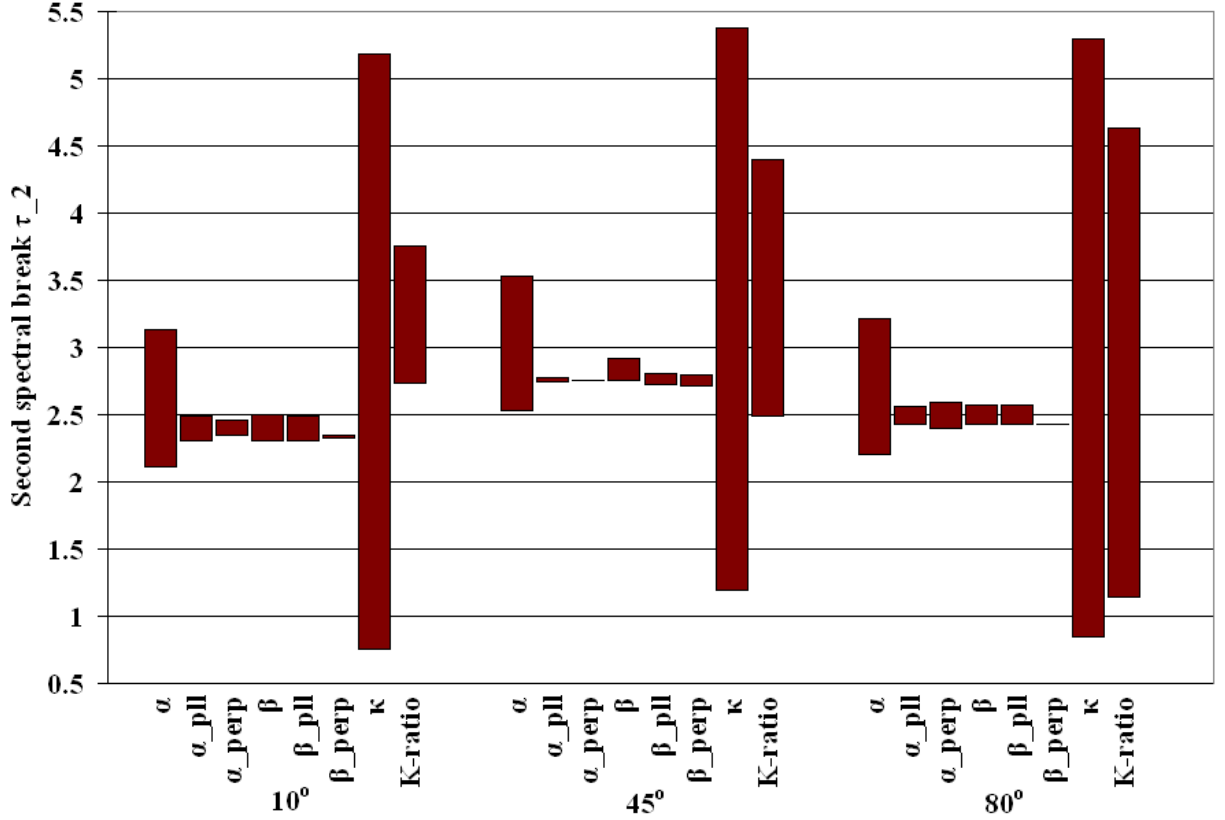


Fig. 28(c).— This figure compares the influences of different magnetic field spectral parameters on the second spectral break $\log_e(\omega/\omega_o\gamma^2) = \tau_2$ of the radiation spectrum. For each spectral parameter (indicated on the bottom axis) the graph indicates the range between the maximum and minimum values of τ_2 obtained by our variations of that parameter (The parameter variations are as indicated in the previous figure and described in detail in section 5.)

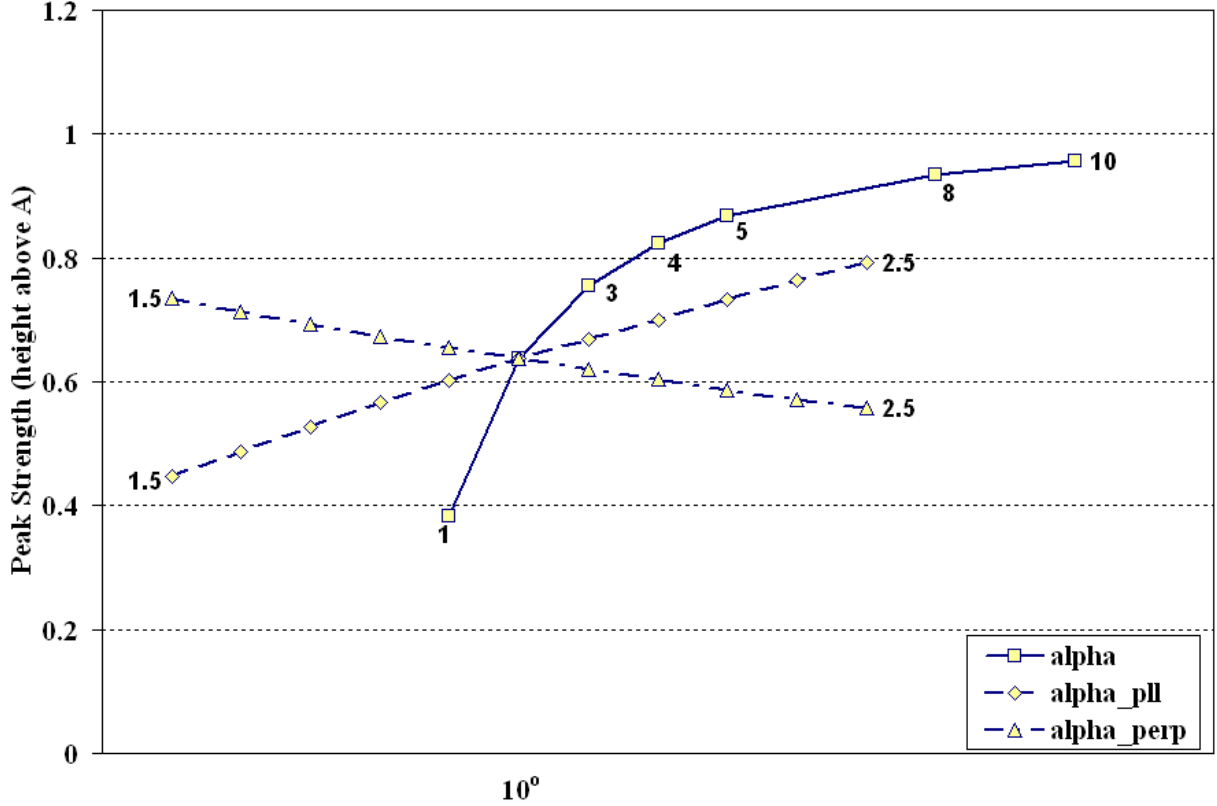


Fig. 29(a).— This figure shows the variation of peak strength, namely the peak height relative to the initial amplitude, for changes in the magnetic field parameters α_i . We present results only for $\theta = 10^\circ$ because at $\theta = 45^\circ$ and $\theta = 80^\circ$ the spectrum is unpeaked. The solid line with square data points indicates the behavior when α is jointly varied (over a range from 1 to 10, as indicated) in both the transverse and parallel magnetic field equation: $\alpha = \alpha_{\perp} = \alpha_{\parallel}$. The dotted lines show the effect of varying α_{\perp} (triangular data points) and α_{\parallel} (diamond data points) individually, from 1.5 to 2.5, in increments of 0.1.

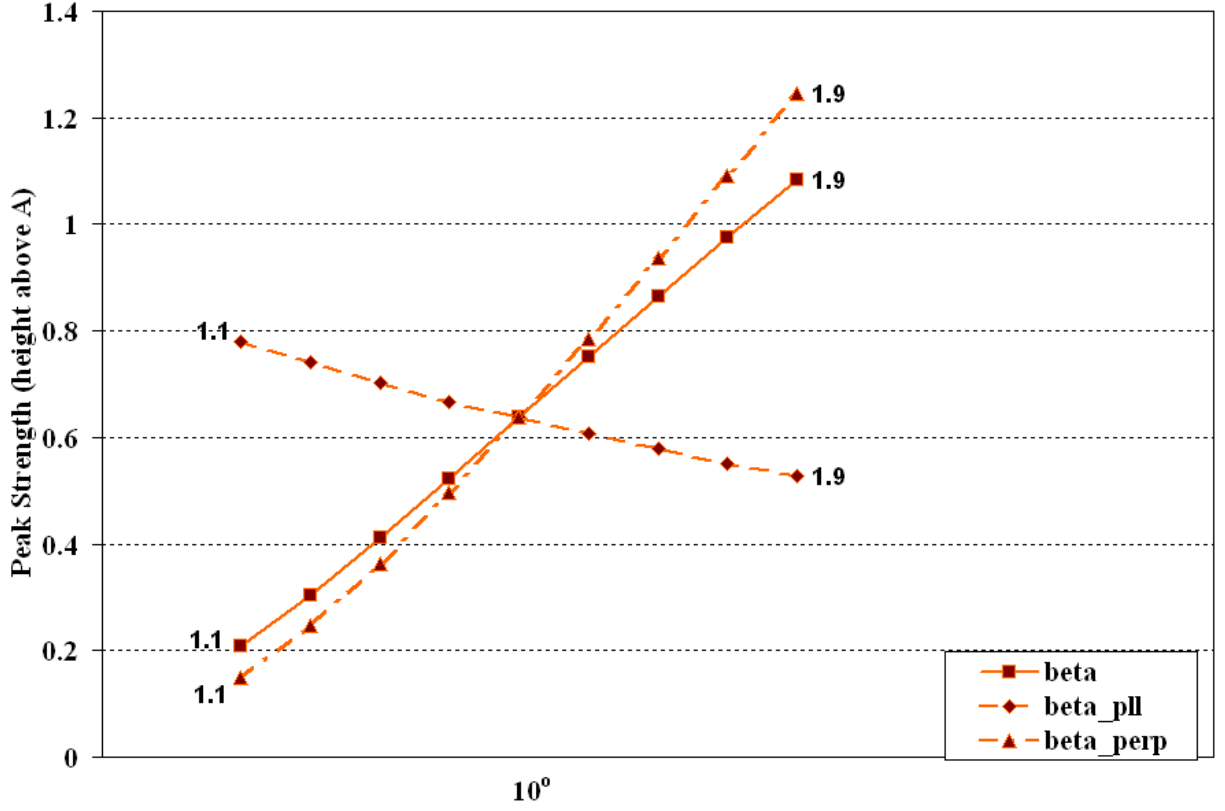


Fig. 29(b).— This figure shows the variation of peak strength (the peak height relative to the initial amplitude) with changes in the magnetic field parameters β_i . The solid line with square data points indicates the behavior when β is jointly varied from 1.1 to 1.9, in increments of 0.1, in both the transverse and parallel magnetic field equation: $\beta = \beta_{\perp} = \beta_{\parallel}$. The dotted lines show the effect of varying β_{\perp} (triangular data points) and β_{\parallel} (diamond data points) individually, also from 1.1 to 1.9, in increments of 0.1.

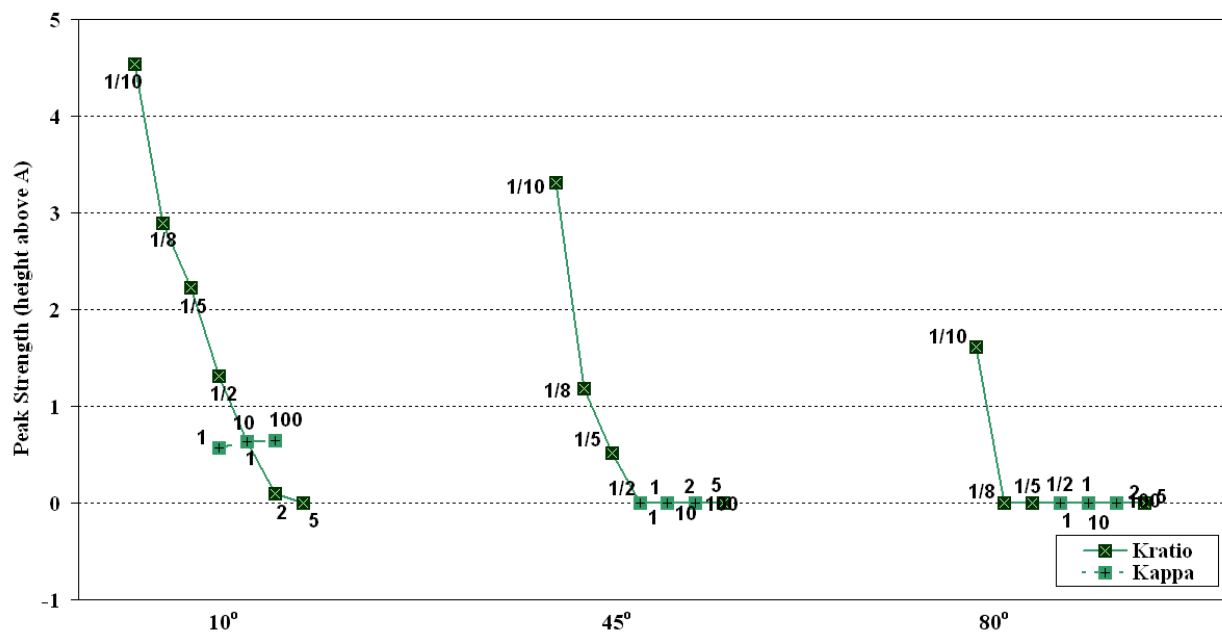


Fig. 29(c).— This figure shows the variation of peak strength (the peak height relative to the initial amplitude) with changes in the magnetic field parameters κ . We vary κ jointly ($\kappa = \kappa_{\perp} = \kappa_{\parallel}$) by powers of 10, from 1 to 100, as indicated. We also vary κ_{\perp} and κ_{\parallel} relative to one another by changing the ratio $K = \kappa_{\perp}/\kappa_{\parallel}$ through a range of values as indicated. Note that we obtain peaked forms of the spectra for larger θ when we have small values of the ratio K .